

Réunions AstroDoc

Le travail de l'équipe DJIN 07/12/2017



Equipe DJIN



Les journaux indexés

Priorité 1

A&A	Astronomy & Astrophysics
AJ	Astronomical Journal
ApJ	AstroPhysical Journal
ApJS	AstroPhysical Journal Supplement Series
MNRAS	Monthly Notices of the Royal Astronomical Society
Natur	Nature
PASJ	Publication of the Astronomical Society of Japan
PASP	Publication of the Astronomical Society of the Pacific
Sci	Science

Les journaux indexés

	Références par an (2016)
A&A	1852
AJ	388
ApJ	3758
ApJS	183
MNRAS	3258
Natur	40
PASJ	123
PASP	94
Sci	12

Total : 9708 références en 2016

Les journaux indexés

Priorité 2

Aca	Acta Astronomica
Raa	Research in Astronomy and Astrophysic
Atel	Astronomers Telegram
GCN	Gamma-ray Burst Coordinates Network
NewA	New Astronomy

Les journaux indexés

Priorité 2

Références par année (2016)

Aca	16 (2015)
Raa	103
Atel	1626 (2015)
GCN	1500
NewA	69

Total : 3314 références

Les journaux indexés

Total : 13022 références par an

Soit 18 articles par jour

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Dispatcher les références aux autres équipes (Dictionnaire, VizieR, Cosim) et aux astronomes pour expertise scientifique et mise à jour des données fondamentales

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Des spécificités à chaque publication (objets taggés et =g0= pour A&A, les données à mettre en =e= dans l'ApJS,...)

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Des spécificités à chaque publication (objets taggés et =g0= pour A&A, ...)

En raison de notre place dans la chaîne documentaire, nous sommes amenés à constater et signaler des erreurs/problèmes, dans Simbad, dans la mise à jour, quand on les remarque mais n'oubliez pas :



18 articles par jour

La notice bibliographique « BIBCODE »

DJIN, mais pas seulement

- * Récupérée de l'Editeur (A&A, ApJ, AJ, MNRAS, PASP,...)
- * Créée via DJIN1
- * Créée manuellement (PASJ, Atel,...)
Programme pour le faire de manière automatique
en cours

TRAITER UN ARTICLE Sans DJIN

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2017A&A...606A..26W: WEDEMEYER S., KUCINSKAS A., KLEVAS J. and LUDWIG H.-G.
<Astron. Astrophys., 606A, 26-26 (2017)>
Three-dimensional hydrodynamical CO^5^BOLD model atmospheres of red giant stars. VI. First chromosphere model of a
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annot.pdf — Three-dimensional hydrodynamical CO5BOLD model atmospheres of red giant stars - VI. First chromosphere model of a late-type giant

Fichier Edition Affichage Aller à Signets Aide

12 sur 12 125%

(An existing (or the modelled star) as a natural upper boundary of a chromosphere. Despite the simplifications, the first model presented here already exhibits a very intermittent and dynamic chromosphere, which is in line with chromospheres modelled for other stellar types. The presented synthetic intensity maps for spectral lines and continua formed in the chromosphere clearly show that the exhibited spatial and temporal variations cannot be correctly described with a static 1D model atmosphere. While 1D models have become very elaborate (see e.g. Dupree et al. 2016), the continued development of adequate 3D (magneto-) hydrodynamical model atmospheres is ultimately needed.

Acknowledgements. We thank V. Dobrovolskas for providing the VUES spectrum of [Aldebaran](#) prior to its publication. This work was supported by a grant from the Research Council of Lithuania (MIP-089/2015). S.W. acknowledges support by a grant from the Research Council of Lithuania (VIZ-TYR-158).

References

Aurière, M., Konstantinova-Antova, R., Charbonnel, C., et al. 2015, *A&A*, 574, A90

Avrett, E. H., & Loeser, R. 2008, *ApJS*, 175, 229

Ayres, T. R., & Testerman, L. 1981, *ApJ*, 245, 1124

Bagnulo, S., Jehin, E., Ledoux, C., et al. 2003, *Messenger*, 114, 10

Bergemann, M., Serenelli, A., Schönrich, R., et al. 2016, *A&A*, 594, A120

Blackwell, D. E., Lynas-Gray, A. E., & Peiford, A. D. 1991, *A&A*, 245, 567

Blanco-Caesares, S., Soubiran, C., Jofré, P., & Heiter, U. 2014, *A&A*, 566, A98

Caffau, E., Ludwig, H.-G., Steffen, M., et al. 2011, *Sol. Phys.*, 268, 255

Carlsson, M. 1986, 'A Computer Program for Solving Multi-Level Non-LTE Radiative Transfer Problems in Moving or Static Atmospheres, Uppsala Astronomical Observatory: Report No. 33

Gustafsson, B., Edvardsson, B., Eriksson, K., et al. 2008, *A&A*, 486, 951

Harper, G. M., Brown, A., & Redfield, S. 2011, in 16th Cambridge Workshop on Cool Stars, Stellar Systems, and the Sun, eds. C. Johns-Krull, M. K. Browning, & A. A. West, *ASP Conf. Ser.*, 448, 1145

Harper, G. M., O'Riain, N., & Ayres, T. R. 2013, *MNRAS*, 428, 2064

Judge, P. G., & Carpenter, K. G. 1998, *ApJ*, 494, 828

Jurgenson, C., Fischer, D., McCracken, T., et al. 2016, *J. Astron. Instr.*, 5, 1650003

Liseau, R., Vlemmings, W., Bayo, A., et al. 2015, *A&A*, 573, L4

Ludwig, H.-G., Jordan, S., & Steffen, M. 1994, *A&A*, 284, 105

Ludwig, H.-G., Caffau, E., Steffen, M., et al. 2009, *Mem. Soc. Astron. It.*, 80, 711

McMurry, A. D. 1999, *MNRAS*, 302, 37

McMurry, A. D., & Jordan, C. 2000, *MNRAS*, 313, 423

McMurry, A. D., Jordan, C., & Carpenter, K. G. 1999, *MNRAS*, 302, 48

Meszáros, S., Dupree, A., & Szentgyörgyi, A. 2008, *AJ*, 135, 1117

Narain, U., & Ulmschneider, P. 1996, *Space Sci. Rev.*, 75, 453

Peter, H., & Judge, P. G. 1999, *ApJ*, 522, 1148

Richichi, A., Dyachenko, V., Pandey, A. K., et al. 2017, *MNRAS*, 464, 231

Robinson, R. D., Carpenter, K. G., & Brown, A. 1998, *ApJ*, 503, 396

Rutten, R. J. 2007, in *The Physics of Chromospheric Plasmas*, eds. P. Heinzel, I. Dorotović, & R. J. Rutten, *ASP Conf. Ser.*, 368, 27

Skartlien, R., Stein, R. F., & Nordlund, Å. 2000, *ApJ*, 541, 468

Steffen, M., Prakapavičius, D., Caffau, E., et al. 2015, *A&A*, 583, A57

Vecchio, A., Cauzzi, G., & Reardon, K. P. 2009, *A&A*, 494, 269

Vernazza, J. E., Avrett, E. H., & Loeser, R. 1981, *ApJS*, 45, 635

Vieytes, M., Maas, P., Cacciari, C., Origlia, L., & Pancino, E. 2011, *A&A*, 526, A4

Vlemmings, W. H. T., Ramstedt, S., O'Gorman, E., et al. 2015, *A&A*, 577, L4

Wedemeyer, S., Freytag, B., Steffen, M., Ludwig, H.-G., & Holweger, H. 2004, *A&A*, 414, 1121

Wedemeyer, S., Ludwig, H.-G., & Steiner, O. 2013, *Astron. Nachr.*, 334, 137

Wedemeyer, S., Bastian, T., Brajša, R., et al. 2016, *Space Sci. Rev.*, 200, 1

Wedemeyer-Böhm, S., Lagg, A., & Nordlund, Å. 2009, *Space Sci. Rev.*, 144, 317

Wedemeyer-Böhm, S., Scullion, E., Steiner, O., et al. 2012, *Nature*, 486, 505

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2017A&A...606A..26W: WEDEMEYER S., KUCINSKA A., KLEVAS J. and LUDWIG H.-G.
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TRAITER UN ARTICLE Sans DJIN

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2017A&A...606A..26W: WEDEMEYER S., KUCINSKAS A., KLEVAS J. and LUDWIG H.-G.
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TRAITER UN ARTICLE

Avec DJIN

1) Ouvrir avec DJIN

2) Eliminer le bruit

3) Lecture de l'article en diagonale

4) Identification des objets restants

5) Verify

6) Simulation

7) Exécution



Ouvrir avec DJIN

The screenshot shows the DJIN software interface. At the top, the window title is "DJIN - 2017A&A...605A..91D". Below the title bar is a menu bar with "File", "Name", "Identifier", "Search", "Configuration", and "Help". A toolbar contains various icons for file operations and search. The main interface is divided into two panes. The left pane, titled "12 object names (60)", lists several astronomical objects with their identifiers and counts in parentheses: #17 (1), #31 (2), 0004-6361 (1), 1400-1800 (1), A4 (1), ESO 8.2 (1), G0 V (4), HD 189733A (3), HD 209458 (31), K1 V (2), Keck-1 (5), and Osiris (8). The right pane displays the details for the selected object, HD 209458. It shows the DOI: 10.1051/0004-6361/201730901 and the title: "Astronomy_ Astrophysics Spatially resolved spectroscopy across stellar surfaces II. High-resolution spectra across HD 209458 (G0 V)". Below the title, there is a "Text" section with the abstract of the paper, starting with "Dainis Dravins, Hans-Günter Ludwig, Erik Dahlén, and Hiva Pazira 1. Spatially resolved stellar spectra Three-dimensional and time-dependent hydrodynamic simulations provide realistic descriptions of the atmospheres of various classes of stars, and spectra computed from such models can be used to determine precise properties of the star and its exoplanets. To constrain and evolve such models, observations beyond the ordinary spectrum of integrated starlight are desirable. Spectral-line syntheses in 3D atmospheres show a rich variety of phenomena characterizing stellar hydrodynamics, which are seen especially in the gradual changes of photospheric line profile strengths, shapes, asymmetries, and wavelength shifts from the center of the disk toward the stellar limb. However, direct comparisons between theory and spectral-line observations have in the past only been possible for the spatially resolved Sun (e.g., Lind et al. 2017). In Paper I (Dravins et al. 2017), we examined theoretically predicted spatially resolved signatures for a group of main-sequence stellar models with temperatures between 6730 and 3960 K. Corresponding observations are feasible during exoplanet transits when small stellar surface portions successively become hidden, and differential spectroscopy between different transit phases can provide spectra of small surface segments temporarily hidden behind the planet. The observational requirements were elaborated in Paper I, in which ob-

Finished



1) Ouvrir avec DJIN

2) *Eliminer le bruit*

3) Lecture de l'article en diagonale

4) Identification des objets restants

5) Verify

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7) Exécution



Éliminer le bruit

DJIN - 2017A&A...605A..91D

File Name Identifier Search Configuration Help

Journal : A&A Volume : 605 Bibcode : 2017A&A...605A..91D

12 object names (60)

- #17 (1)
- #31 (2)
- 0004-6361 (1)
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- A4 (1)
- ESO 8.2 (1)
- G0 V (4)
- HD 189733A (3)
- HD 209458 (31)
- K1 V (2)
- Keck-1 (5)
- Osiris (8)

Text

DOI: 10.1051/0004-6361/201730901

Astronomy_ Astrophysics Spatially resolved spectroscopy across stellar surfaces II. High-resolution spectra across **HD 209458 (**G0 V**)**

Text

Devin Dressing, Hans-Günter Ludwig, Erik Dehnen, and Ugo Perina 1. Spatially resolved stellar

For HD 209458, hundreds of archive spectra were retrieved from several different observatories and examined for their suitability. In particular, the ESO Science Archive Facility was examined for data from the UVES and HARPS spectrometers at the VLT Kueyen on Paranal and the 3.6 m telescope on La Silla, respectively; the Keck Observatory Archive for data from its High Resolution Echelle Spectrometer (HIRES) at **Keck-1** on Maunakea; and the SMOKA science archive (Subaru Mitaka Okayama Kiso Archive); for spectra from the High Dispersion Spectrograph (HDS) at the Subaru telescope, also on Maunakea.

and evolve such models, observations beyond the ordinary spectrum of integrated starlight are desirable.

Spectral-line syntheses in 3D atmospheres show a rich variety of phenomena characterizing stellar hydrodynamics, which are seen especially in the gradual changes of photospheric line profile strengths, shapes, asymmetries, and wavelength shifts from the center of the disk toward the stellar limb. However, direct comparisons between theory and spectral-line observations have in the past only been possible for the spatially resolved Sun (e.g., Lind et al. 2017). In Paper I (Dravins et al. 2017), we examined theoretically predicted spatially resolved signatures for a group of main-sequence stellar models with temperatures between 6730 and 3960 K. Corresponding observations are feasible during exoplanet transits when small stellar surface portions successively become hidden, and differential spectroscopy between different transit phases can provide spectra of small surface segments temporarily hidden behind the planet. The observational requirements were elaborated in Paper I, in which ob-

In this Paper II,

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Éliminer le bruit

DJIN - 2017A&A...605A..91D

File Name Identifier Search Configuration Help

Journal : A&A Volume : 605 Bibcode : 2017A&A...605A..91D

12 object names (60)

- #17 (1)
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- ESO 8.2 (1)

Text

DOI: 10.1051/[0004-6361/201730901](https://doi.org/10.1051/0004-6361/201730901)

Astronomy_ Astrophysics Spatially resolved spectroscopy across stellar surfaces II. High-

Snellen, A. Collier Cameron, and K. Horne. At Lund Observatory, contributions to the examination of archival spectra from different observatories were made also by Tiphaine Lagadec and Joel Wallenius. H.G.L. acknowledges financial support by the Sonderforschungsbereich SFB881 "The Milky Way System" (subproject A4) of the German Research Foundation (DFG). The work by D.D. was performed in part at the Aspen Center for Physics, which is supported by National Science Foundation grant PHY-1066293. D.D. also acknowledges stimulating stays as a Scientific Visitor at the European Southern Observatory in Santiago de Chile. We thank the referee for constructive and valuable comments.

Osiris (8)

Dainis Dravins, Hans-Günter Ludwig, Erik Dahlén, and Hiva Pazira 1. Spatially resolved stellar spectra Three-dimensional and time-dependent hydrodynamic simulations provide realistic descriptions of the atmospheres of various classes of stars, and spectra computed from such models can be used to determine precise properties of the star and its exoplanets. To constrain and evolve such models, observations beyond the ordinary spectrum of integrated starlight are desirable. Spectral-line syntheses in 3D atmospheres show a rich variety of phenomena characterizing stellar hydrodynamics, which are seen especially in the gradual changes of photospheric line profile strengths, shapes, asymmetries, and wavelength shifts from the center of the disk toward the stellar limb. However, direct comparisons between theory and spectral-line observations have in the past only been possible for the spatially resolved Sun (e.g., Lind et al. 2017). In Paper I (Dravins et al. 2017), we examined theoretically predicted spatially resolved signatures for a group of main-sequence stellar models with temperatures between 6730 and 3960 K. Corresponding observations are feasible during exoplanet transits when small stellar surface portions successively become hidden, and differential spectroscopy between different transit phases can provide spectra of small surface segments temporarily hidden behind the planet. The observational requirements were elaborated in Paper I, in which ob-

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Éliminer le bruit

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File Name Identifier Search Configuration Help

Journal: A&A Volume: 605 Bibcode: 2017A&A...605A..91D

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Astronomy - Astrophysics Spatially Resolved Spectroscopy across stellar surfaces II. High-resolution spectra across HD 209458 (G0 V)

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[KBS98] -1, [VFK98] -1

Fig. 7. Wavelength positions of selected individual Fe I lines with the exposure-averaged wavelengths shifted to those of the measured Rossiter-McLaughlin effect (dashed gray). The increased noise around exposure #17 is caused by the S/N then dropping to "only" 300; cf.



Éliminer le bruit : résultat

The screenshot shows a web browser window with the following details:

- Browser tab: DJIN - 2017A&A...605A..91D
- Page title: File Name Identifier Search Configuration Help
- Journal: A&A
- Volume: 605
- Bibcode: 2017A&A...605A..91D
- Left sidebar: 3 object names (42)
 - HD 189733A (3)
 - HD 209458 (31)
 - Osiris (8)
- Main content area:
 - Text: DOI: 10.1051/0004-6361/201730901
 - Section header: **Astronomy_ Astrophysics Spatially resolved spectroscopy across stellar surfaces II. High-resolution spectra across HD 209458 (G0 V)**
 - Text: Dainis Dravins, Hans-Günter Ludwig, Erik Dahlén, and Hiva Pazira 1. Spatially resolved stellar spectra Three-dimensional and time-dependent hydrodynamic simulations provide realistic descriptions of the atmospheres of various classes of stars, and spectra computed from such models can be used to determine precise properties of the star and its exoplanets. To constrain and evolve such models, observations beyond the ordinary spectrum of integrated starlight are desirable. Spectral-line syntheses in 3D atmospheres show a rich variety of phenomena characterizing stellar hydrodynamics, which are seen especially in the gradual changes of photospheric line profile strengths, shapes, asymmetries, and wavelength shifts from the center of the disk toward the stellar limb. However, direct comparisons between theory and spectral-line observations have in the past only been possible for the spatially resolved Sun (e.g., Lind et al. 2017). In Paper I (Dravins et al. 2017), we examined theoretically predicted spatially resolved signatures for a group of main-sequence stellar models with temperatures between 6730 and 3960 K. Corresponding observations are feasible during exoplanet transits when small stellar surface portions successively become hidden, and differential spectroscopy between different transit phases can provide spectra of small surface segments temporarily hidden behind the planet. The observational requirements were elaborated in Paper I, in which ob-

At the bottom of the browser window, there is a status bar with the text: [KBS98] -1, [VFK98] -1



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3) *Lecture de l'article en diagonale*

Repérage des -oublis/mauvaises détections de DJIN

-tables

-données disponibles en ligne

-données à ajouter aux objets

4) Identification des objets restants

5) Verify

6) Simulation

7) Exécution

Lecture de l'article en diagonal

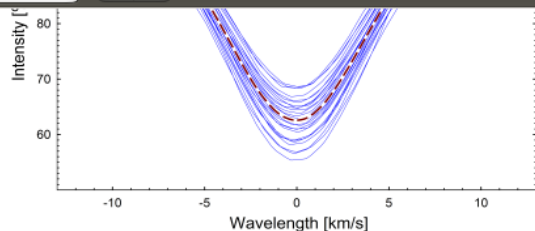


Fig. 8. Averaging 26 “similar” photospheric Fe I lines in HD 209458, selected to be largely unblended and of closely similar strengths (Fig. 5). Their average (dashed) “synthesizes” a representative profile of such a strength in that particular wavelength region. For this reference profile averaged over several exposures outside transit, the photometric signal-to-noise ratio approaches ~ 7000 .

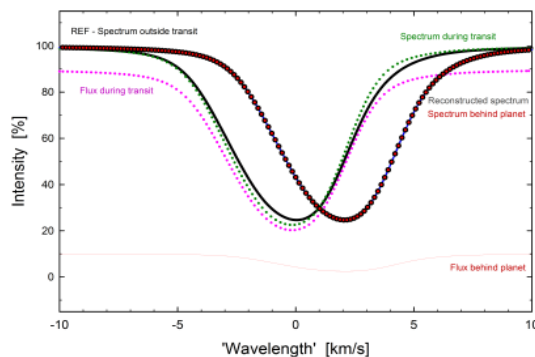


Fig. 9. Different spectral-line components treated in the reconstruction of spatially resolved spectra. The spectrum behind the planet is obtained as that line profile (weighted with the amount of flux temporarily obscured by the planet) that – summed with the temporarily observed line profile – produces the stellar reference profile outside of transit. For clarity, the planetary signal here is greatly exaggerated.

HD 209458 by Hayek et al. (2012) from 3D stellar model atmospheres in passbands including the red SDSS r' , whose effective wavelength of 620.4 nm (Fukugita et al. 1996) closely coincides with the average for our spectral-line selections. However, we stress that our method of spatially resolved line reconstruction does not depend on any theoretical predictions of limb darkening; only the product of planetary area and stellar local brightness is required and such data could be taken directly from observed transit photometry. Separating the values for the planet area and limb darkening makes it easier to discuss error budgets, and a limb-darkening model enables us to extrapolate continuum intensities to the stellar disk center that is not sampled during the exoplanet transit.

6.2. Reconstructing spatially resolved profiles

The principle for profile reconstruction is shown in Fig. 9. The spectroscopically observed profiles are different from each

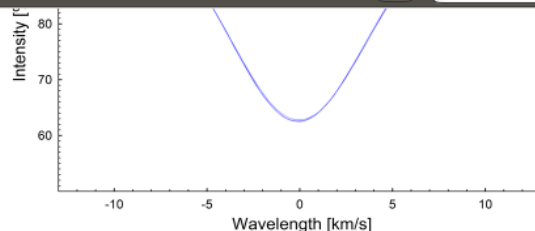


Fig. 10. Representative profiles of the weaker Fe I line and its gradual changes. Each profile, which is the average of 26 Fe I lines, is plotted for 14 successive exposures during the exoplanet transit, showing the degree of uniformity reached. To appreciate the slight changes during the transit, this plot should be viewed highly magnified; corresponding line ratios are in Fig. 11. The photometric signal-to-noise ratio here is ~ 2500 .

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To compute the flux temporarily hidden during transit, the projected planet area was taken as 1.5% of the stellar disk, with limb darkening from Hayek et al. (2012) for the passband SDSS r' ; these values are deduced from photometry. The center-to-limb position of the planet from the transit trajectory was computed from the impact parameter $= 0.51$, as deduced by various authors from the measured Rossiter-McLaughlin effect.



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1) Ouvrir avec DJIN

2) Eliminer le bruit

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4) Identification des objets restants

5) Verify

6) Simulation

7) Exécution



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Différents outils -références citées dans l'article

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- 3 object names (42)
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2. Candidate stars with transiting planets

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Fichier Edition Graphic
Planet orbiting around \object{HD 209458}, see details about \exoplanet{HD 209458}{b} in the \exoEncyclopedie.

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#2 | HD 189733b | HD 189733b |PI |20 00 43.71347 +22 42 39.0645|822

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Coord FK5(1950): 19 58 34.161+22 34 31.79 (Opt ) [0.176 0.138 0] A 2016A&A...595A...2G
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sgal= 323.3058+72.9970
ecl= 308.0478974+42.1842561 [A] 2016A&A...595A...2G
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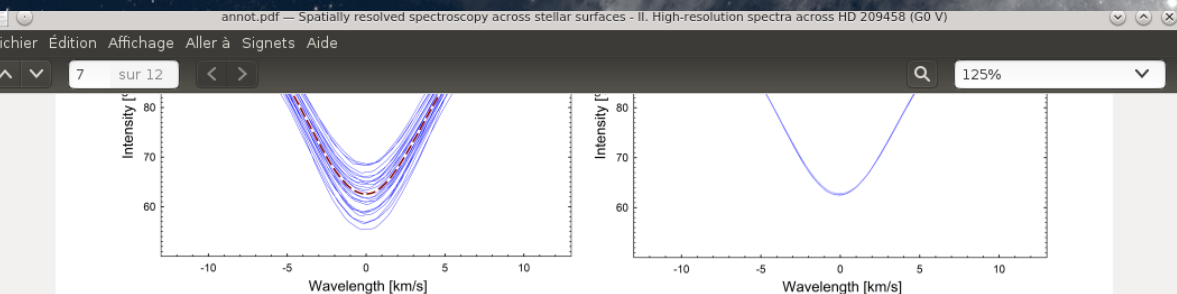


Fig. 8. Averaging 26 “similar” photospheric Fe I lines in HD 209458, selected to be largely unblended and of closely similar strengths (Fig. 5). Their average (dashed) “synthesizes” a representative profile of such a strength in that particular wavelength region. For this reference profile averaged over several exposures outside transit, the photometric signal-to-noise ratio approaches ~7000.

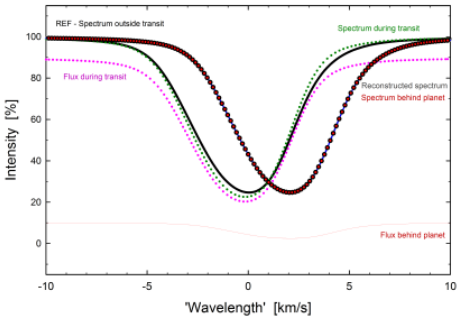


Fig. 9. Different spectral-line components treated in the reconstruction of spatially resolved spectra. The spectrum behind the planet is obtained as that line profile (weighted with the amount of flux temporarily obscured by the planet) that – summed with the temporarily observed line profile – produces the stellar reference profile outside of transit. For clarity, the planetary signal here is greatly exaggerated.

HD 209458 by Hayek et al. (2012) from 3D stellar model atmospheres in passbands including the red SDSS r' , whose effective wavelength of 620.4 nm (Fukugita et al. 1996) closely coincides with the average for our spectral-line selections. However, we stress that our method of spatially resolved line reconstruction does not depend on any theoretical predictions of limb darkening; only the product of planetary area and stellar local brightness is required and such data could be taken directly from observed transit photometry. Separating the values for the planet area and limb darkening makes it easier to discuss error budgets, and a limb-darkening model enables us to extrapolate continuum intensities to the stellar disk center that is not sampled during the exoplanet transit.

6.2. Reconstructing spatially resolved profiles

The principle for profile reconstruction is shown in Fig. 9. The spectroscopically observed profiles are different from each

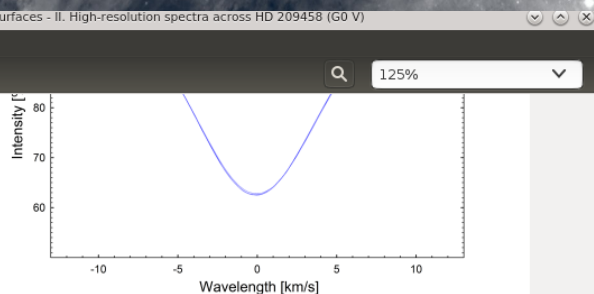


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Identifier Search Configuration Help

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- 4) Identification des objets restants
- 5) *Verify***
- 6) Simulation
- 7) Exécution



TRAITER UN ARTICLE

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Permet de visualiser le script et ses éventuels bugs
- 7) Exécution

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 - Osiris (8)

Simulation

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HD 189733A

The screenshot shows a software application window titled "DJIN - 2017A&A...605A..91D". The interface includes a menu bar with "File", "Name", "Identifier", "Search", "Configuration", and "Help". Below the menu bar, there are fields for "Journal : A&A", "Volume : 605", and "Bibcode : 2017A&A...605A..91D".

On the left side, there is a tree view showing a hierarchy of object names: "3 object names (42)", "HD 189733A (3)", "Text (2)", "Caption (1)", "HD 209458 (31)", and "Osiris (8)".

The main area displays a list of identifiers: "HD 189733", "HD 209458", and "NAME Osiris". A context menu is open over this list, containing buttons for "Verify", "Visualize...", "Simulate", "Execute...", "Open script...", "Delete...", "Commands...", "VizieR...", "Coordinates...", "Properties...", "Duplicate", "Find/replace...", and "Close".

An "Execution" dialog box is overlaid on the context menu, asking "Do you really want to run the script in SIMBAD ?" with "OK" and "Annuler" buttons.

At the bottom of the main window, there are checkboxes for "Status processed", "Question", and "Raw idents", along with a "Bibcode" icon. The status bar at the very bottom displays "HD 189733A".

Background text from a document is visible on the right side of the window, mentioning "Sun (e.g., Lind et al.)", "spatially resolved", "between 6730 and", "small stellar", "different", "hidden behind the", "In this Paper II,", "existence of", "As a", "middle of the", "noise are required;", "extraction of", "positions across", "hydrodynamic models", "exhaustive future", "exoplanets cover", "low noise permit the", "transiting planets for di", "only Jupiter-size (or", "requirements limit these", "studies to the brightest host stars with the largest planets, in particular HD 209458 (G0 V) and HD 189733A (K1 V), and these were selected for detailed investigations. The first of these is the".

Journal : A&A

- 3 object names (42)
 - HD 189733A (3)
 - Text (2)
 - 1
 - 2
 - Caption (1)
 - HD 209458 (31)
 - Osiris (8)

Execution

Flux:

Mesurements:

Number of References (1200):

Notes:

(S) 15-Jan-2006 - fo ref:~ obj:HD 209458b
 Planet orbiting around \object{HD 209458}, see details about \exoplanet{HD 209458}{b} in the \exoEncyclopedie.

NAME Osiris : update > v

2017A&A...605A..91D : update OBJ > .
 2017A&A...605A..91D : update > s ct
 Commentaire privé modifié en :
 ~
 2017A&A...605A..91D : update > v

O[B] | B[IB] | h[elp] : update > q

Statistiques :
 {Attention ! Cette reference etait deja la =3}

*** BYE !! fin de Simup ***

Close

HD 189733A

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Spatially resolved spectroscopy across stellar surfaces. II. High-resolution spectra across HD 209458 (G0 V).

DRAVINS D.; LUDWIG H.-G.; DAHLEN E.; PAZIRA H.

Abstract (from CDS): Context. High-resolution spectroscopy across spatially resolved stellar surfaces aims at obtaining spectral-line profiles that are free from rotational broadening; the gradual changes of these profiles from disk center toward the stellar limb reveal properties of atmospheric fine structure, which are possible to model with 3D hydrodynamics. Aims. Previous such studies have only been carried out for the Sun but are now extended to other stars. In this work, profiles of photospheric spectral lines are retrieved across the disk of the planet-hosting star HD 209458 (G0 V). Methods. During exoplanet transit, stellar surface portions successively become hidden and differential spectroscopy provides spectra of small surface segments temporarily hidden behind the planet. The method was elaborated in Paper I (Dravins et al., 2017A&A...605A..90D), with observable signatures quantitatively predicted from hydrodynamic simulations. Results. From observations of HD 209458 with spectral resolution $\lambda/\Delta\lambda \sim 80000$, photospheric Fe I line profiles are obtained at several center-to-limb positions, reaching adequately high S/N after averaging over numerous similar lines. Conclusions. Retrieved line profiles are compared to synthetic line profiles. Hydrodynamic 3D models predict, and current observations confirm, that photospheric absorption lines become broader and shallower toward the stellar limb, reflecting that horizontal velocities in stellar granulation are greater than vertical velocities. Additional types of 3D signatures will become observable with the highest resolution spectrometers at large telescopes.


Abstract Copyright: © ESO, 2017 European Southern Observatory (ESO) 2017

Journal keyword(s): stars: atmospheres - techniques: spectroscopic - line: profiles - hydrodynamics - planets and satellites: gaseous planets - stars: solar-type

Simbad objects: 3

▶ Full paper

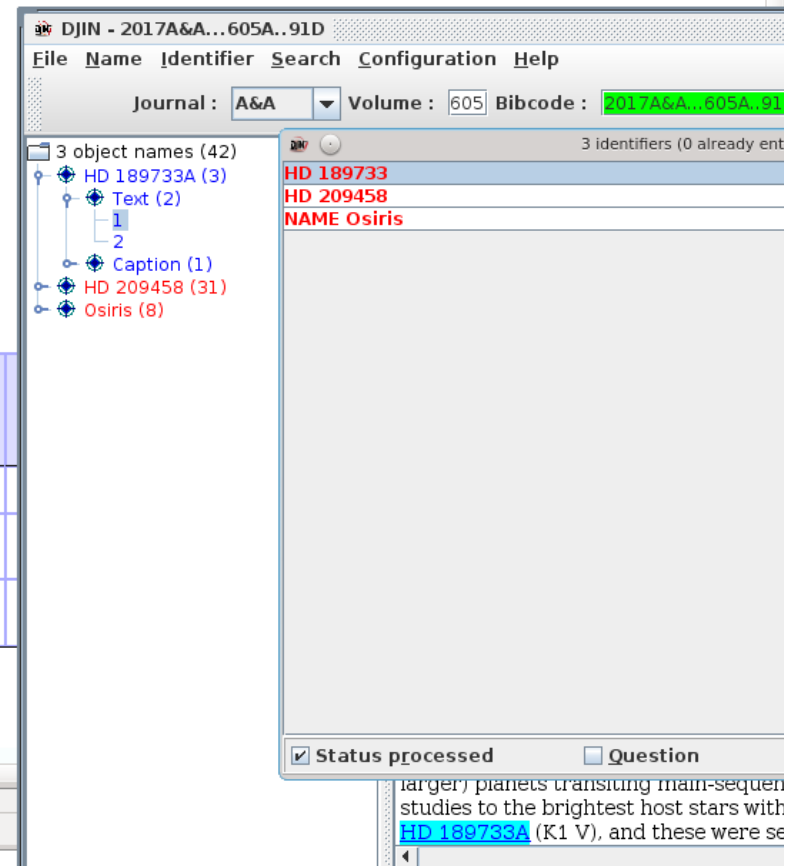
▶ View the reference in ADS

Number of rows : 3 

Show 100 entries

N	Identifier	obj tags	obj count	Otype	ICRS (J2000) RA	ICRS (J2000) DEC	Mag U	Mag B	Mag V
1	HD 189733	cx	3	BY*	20 00 43.7128	+22 42 39.074	9.241	8.578	7.648
2	NAME Osiris	sx	8	Pl	22 03 10.77207	+18 53 03.5430			
3	V* V376 Peg	tascx	31	EP*	22 03 10.7728	+18 53 03.550		8.21	7.63

Showing 1 to 3 of 3 entries



DJIN - 2017A&A...605A..91D

File Name Identifier Search Configuration Help

Journal : A&A Volume : 605 Bibcode : 2017A&A...605A..91D

3 object names (42)

- HD 189733A (3)
 - Text (2)
 - 1
 - 2
 - Caption (1)
- HD 209458 (31)
- Osiris (8)

3 identifiers (0 already entered)

HD 189733

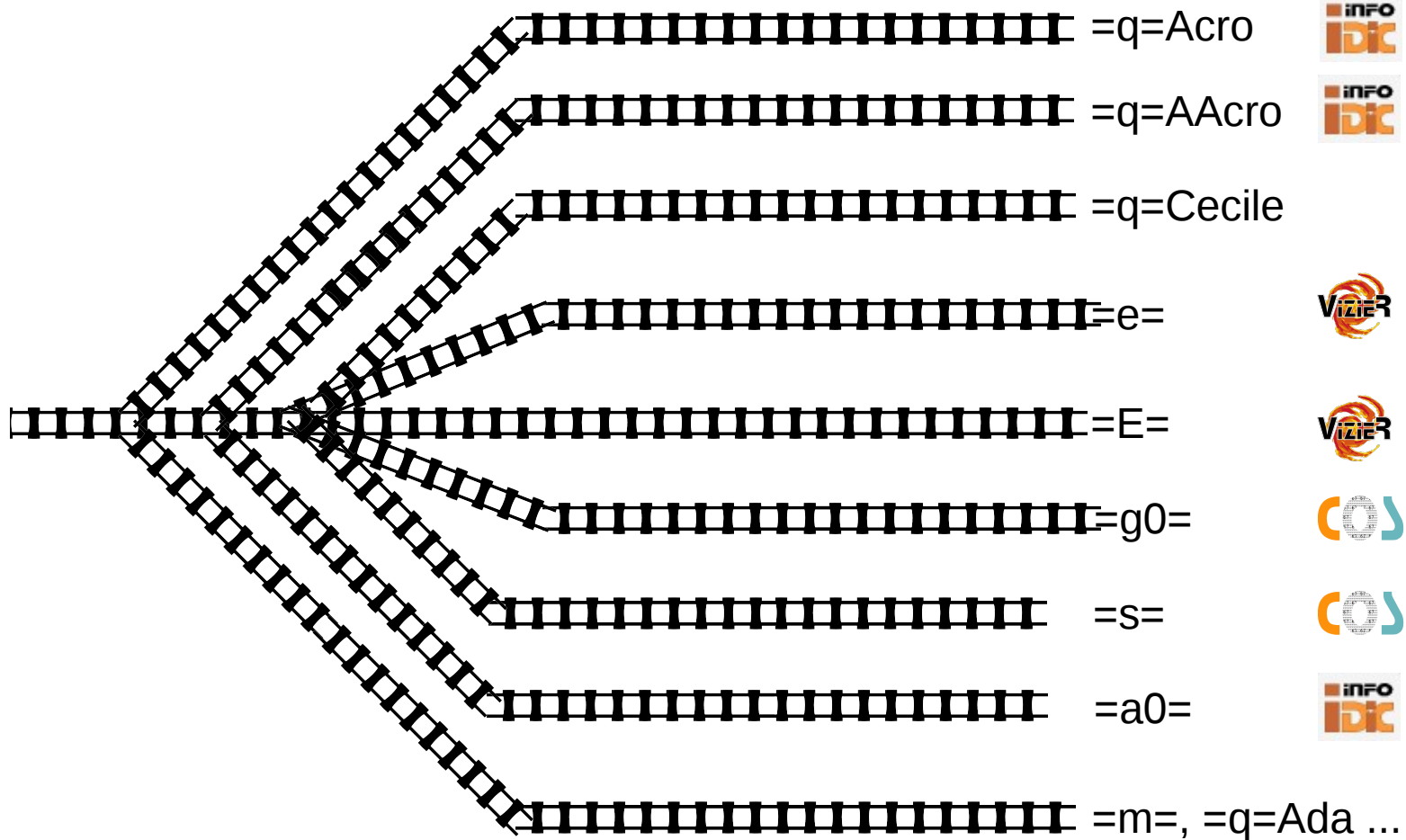
HD 209458

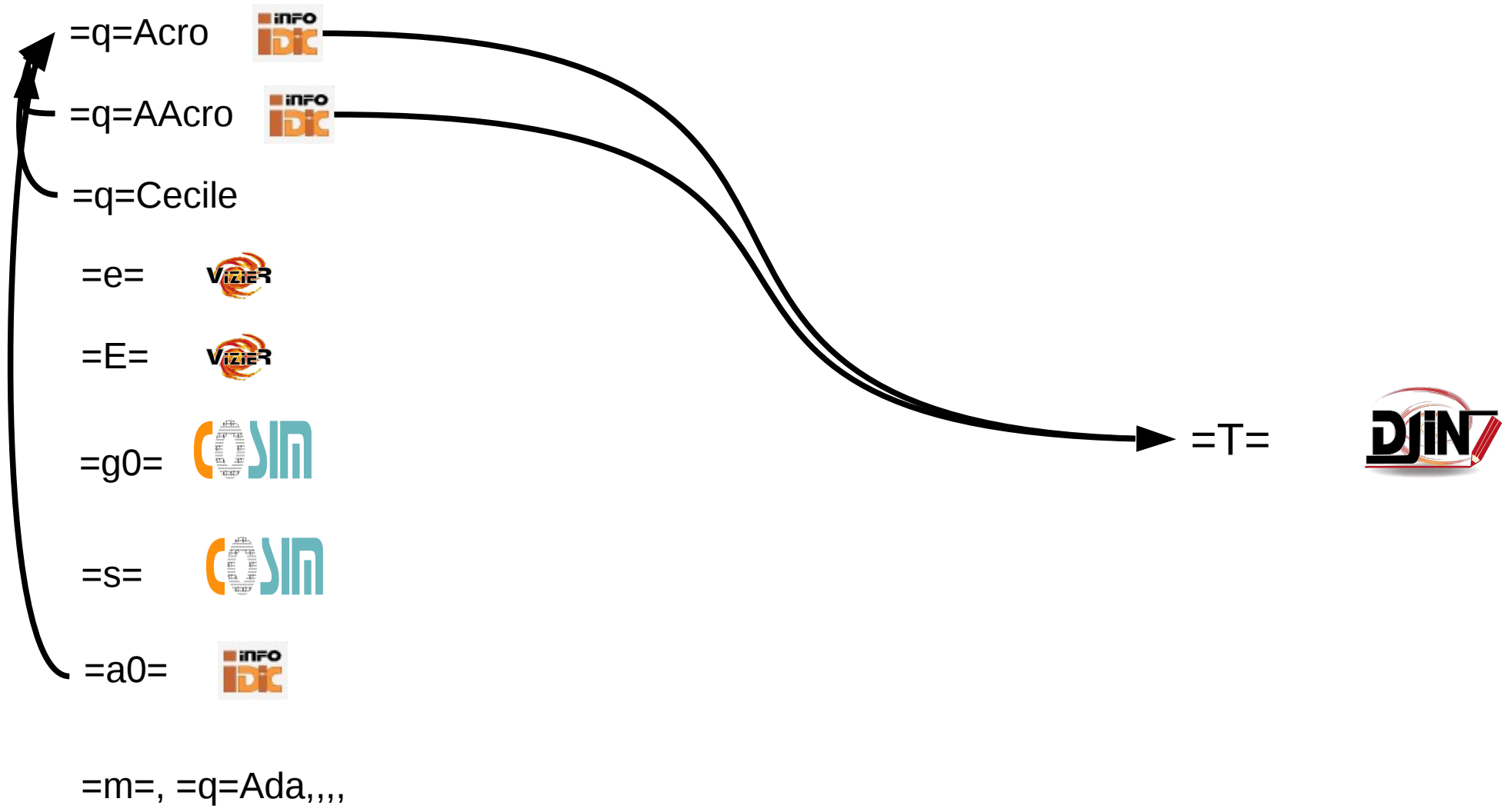
NAME Osiris

Status processed Question

larger) planets transiting main-sequence stars. Studies to the brightest host stars with HD 189733A (K1 V), and these were se

The spectroscopically observed profiles are different from each other from the measured Rossiter-McLaughlin effect.





=q=Acro

Appendix A: Optical identifications

In this Appendix we report the full list of optical counterparts belonging to [A 2142](#) in the same field shown in Fig. 2, and radio/optical overlays both for the extended radio galaxies and for those groups where multiple radio emission has been detected.

Table A.1. Optical identifications with cluster galaxies.

#, GMRT name	α_{J2000} radio	δ_{J2000} radio	$S_{608 \text{ MHz}}$ (mJy)	$S_{234 \text{ MHz}}$ (mJy)	$\log P_{608 \text{ MHz}}$ (W/Hz)	Notes
Optical Catalogue	α_{J2000} opt	δ_{J2000} opt	m_g		z	
<u>#1, GMRT-J 155636+270041</u> WISEPC	15 56 36.99 15 56 37.07	27 00 41.7 27 00 39.7	1.67 15.9	6.52	22.55 0.091117	P
<u>#2, GMRT-J 155642+273324</u> 2MASX	15 56 42.94 15 56 42.96	27 33 24.1 27 33 24.6	3.18 17.0	5.71	22.80 0.088777	P
<u>#3, GMRT-J 155646+270015</u> SDSS	15 56 46.20 15 56 46.25	27 00 15.2 27 00 15.4	0.53 18.9	–	21.99 0.085227	P
<u>#4, GMRT-J 155700+273102</u> 2MASX	15 57 00.56 15 57 00.56	27 31 02.2 27 31 02.7	1.87 17.8	4.42	22.50 0.082248	P
<u>#5, GMRT-J 155703+271812</u> 2MASX	15 57 03.29 15 57 03.34	27 18 12.8 27 18 12.7	1.53 16.8	3.65	22.54 0.094222	P

$$=q=A\text{Acro}$$

also the images in the MOSAIC webpage).

The disappearance of strong IDV in S4 0917+624 after the year 2000 cannot be fully explained with the quenching mechanism via changes of core size in the source, as discussed above. Bernhart et al. (2006) studied the VLBI kinematics of S4 0917+624 with data during 2000–2007 at 5, 15, and 22 GHz, and also did not find clear correlation between the IDV properties and VLBI structure. The disappearance is likely caused by a change of ISM scattering properties, e.g. with the passage of scattering material out of the line of sight to the quasar, as discussed for the fast IDV source J1819+3845, which also ceased at some time in the period between June 2006 and February 2007 (de Bruyn & Macquart 2015), and for the intermittent IDV source PKS 0405–385 (Kedziora-Chudczer 2006).

To identify the possible scattering materials located in the foreground of S4 0917+624, which has the distance ~ 200 pc estimated with the shortest timescale in the appendix by Rickett

=q=Cecile

Table 1. Detected sources.

Source ID	RA	Dec	Size (arcsec × arcsec)	Reference
<i>Map <u>M1</u>/Band 3</i>				
1a	07:47:30.817	-19:17:18.48	1.76'' × 1.57''	USNO <u>B1.0 0707-10151219</u>
2a	07:47:30.520	-19:17:23.39	2.01'' × 1.85''	unknown
Bright source in Fig. 10b	17:47:30.130	-19:17:50.60	2.56'' × 2.42''	<u>2MASX J07473002-1917503</u>
<i>Map <u>M2</u>/Band 7</i>				
1b	07:47:31.206	-19:17:37.38	0.32'' × 0.26''	unknown
2b	07:47:31.642	-19:17:38.85	0.52'' × 0.40''	unknown
3b	07:47:31.228	-19:17:37.35	0.35'' × 0.34''	unknown
1b+3b (as a single source)	07:47:31.220	-19:17:37.35	0.44'' × 0.32''	unknown
4b	07:47:30.982	-19:17:36.53	0.34'' × 0.24''	unknown
<i>Map <u>M3</u>/Band 3 + Band 7</i>				
1c	07:47:31.618	-19:17:35.464	0.39'' × 0.18''	unknown
<u>2c</u>	07:47:30.974	-19:17:35.464	0.41'' × 0.24''	unknown
<u>3c</u>	07:47:31.192	-19:17:37.33	0.57'' × 0.31''	same as 1b+3b in <u>M2</u>

S. Simón-Díaz et al.: New observational clues to understand macroturbulent broadening in massive O- and B-type stars

Table 1. Stars considered for this paper, including information about line broadening, stellar parameters, and the quantity RSk (relative skewness, see Eq. (1)).

Target	SpC	Line	S/N_c	EW	$v \sin i$		v_{mac}	RSk	σ_{RSk}	$\log T_{\text{eff}}$	$\log \mathcal{L}/\mathcal{L}_{\odot}$
					FT	GOF	GOF				
...
<u>HD 16582</u>	<u>B2 IV</u>	Si III	183	138	9	9	19	0.02	0.05	4.34	2.94
<u>HD 17081</u>	<u>B7 V</u>	Mg II	195	294	18	19	24	-0.03	0.02	4.12	2.40
<u>HD 17603</u>	O7.5 Ib(f)	O III	239	317	109	99	115	-0.01	0.02	4.53	4.21
<u>HD 17743</u>	<u>B8 III</u>	Mg II	157	214	48	47	22	0.00	0.06	4.13	2.21
<u>HD 18409</u>	O9.7 Ib	O III	256	179	131	128	<88	0.08	0.16	4.51	3.99
<u>HD 18604</u>	<u>B6 III</u>	Mg II	184	305	131	132	<50	0.07	0.08	4.11	2.44
<u>HD 19820</u>	O8.5 III(n)((f))	O III	316	251	144	147	<54	-0.22	0.11	4.51	3.91
...

Notes. The line used to determine the line-broadening parameters is also indicated, along with its equivalent width and the signal-to-noise ratio of the adjacent continuum. Spectral classifications indicated in Col. 2 must be handled with caution since they come from various sources, not all of which are equally reliable. EW in mÅ, $v \sin i$ and v_{mac} in km s^{-1} , T_{eff} in K. The full table is available at the CDS.

* Full Table 1 is only available at the CDS via anonymous ftp to cdsarc.u-strasbg.fr (130.79.128.5) or via <http://cdsarc.u-strasbg.fr/viz-bin/qcat?J/A+A/597/A22>

¹ Throughout this paper, and following Reed (2003), we use the term OB stars to refer to O- and early-B type stars on the main sequence and their evolved descendants, the B supergiants. The remaining B-type stars (dwarfs and giants) are considered as a separate group.

=e=

Table 1
Data on the *Fermi*/LAT Blazar Candidates (583 Sources)

LAT Name	Counterpart Name	Class	$\log \nu_s$	Γ_γ	$\alpha_{r\gamma}$	α_{rx}	α_{ox}	ρ_s^1	ρ_s^2	ρ_s^3	IT
J2250.3-4206	PMN J2250-4206	BCU I	13.92	2.034	0.750	0.45
J1647.4+4950	SBS 1646+499	BCU I	13.68	2.429	0.759	0.662	1.693	-0.20	-0.38	0.10	Q
J1412.0+5249	SBS 1410+530	BCU I	15.364	2.562	0.875	0.776	2.181	-0.90	-0.88	-0.41	...
J0343.3+3622	OE 367	BCU I	12.28	2.426	0.828	-0.73	Q
J0647.6-6058	PMN J0647-6058	BCU I	13.11	2.337	0.759	-0.07
J0040.5-2339	PMN J0040-2340	BCU I	14.358	1.946	0.768	0.40
J2040.2-7115	PKS 2035-714	BCU I	16.15	1.798	0.795	0.624	...	0.32	0.49
J0151.0-3609	PMN J0151-3605	BCU I	...	2.459	0.772	-0.35
J2336.5+2356	B2 2334+23	BCU I	...	2.134	0.822	0.639	...	-0.24	0.09
J0132.5-0802	PKS 0130-083	BCU I	15.75	1.753	0.807	0.632	1.409	0.28	0.49	-0.31	...
J0618.9-1138	TXS 0616-116	BCU I	13.75	2.470	0.822	-0.75
J1416.0+1325	PKS B1413+135	BCU I	12.865	2.363	0.843	-0.68
J0059.1-5701	PKS 0056-572	BCU I	12.77	2.616	0.827	0.682	1.004	-0.87	-0.64	-0.70	...
J0749.4+1059	TXS 0746+110	BCU I	14.265	2.203	0.795	-0.16
J1604.4-4442	PMN J1604-4441	BCU I	12.947	2.453	0.793	-0.50
J1121.0-0502	PKS 1128-047	BCU I	12.865	2.659	0.825	0.04

= E =

Table 8
List of New Minimum Timings Used for the Analysis

Star	JD Hel.- 2400000	Error (day)	Type	Filter	Source/ Observatory
<u>V773 Cas</u>	48500.8874	0.0095	Prim	Hp	<i>Hipparcos</i>
<u>V773 Cas</u>	54798.48508	0.00154	Prim	BVR	PS—this study
<u>V773 Cas</u>	55062.39534	0.00139	Prim	BVR	PS—this study
<u>V773 Cas</u>	55071.45118	0.00133	Prim	BVR	PS—this study
<u>V773 Cas</u>	55410.39017	0.00197	Prim	BVR	PS—this study
<u>V773 Cas</u>	55419.44651	0.00101	Prim	BVR	PS—this study
<u>V773 Cas</u>	55481.54064	0.00182	Prim	BVR	PS—this study
<u>V773 Cas</u>	55754.50689	0.00038	Sec	I	RU—this study
<u>V773 Cas</u>	55776.49780	0.00026	Prim	R	PS—this study
<u>V773 Cas</u>	55776.49748	0.00066	Prim	I	RU—this study
<u>V773 Cas</u>	55877.40172	0.00047	Prim	R	RU—this study
<u>V773 Cas</u>	56155.54350	0.00069	Sec	R	RU—this study
<u>V773 Cas</u>	56159.42162	0.00105	Prim	R	RU—this study
<u>V773 Cas</u>	56230.57514	0.00062	Sec	C	RU—this study
<u>V773 Cas</u>	56252.56717	0.00075	Prim	C	RU—this study
<u>V773 Cas</u>	56304.31383	0.00040	Prim	C	RU—this study
<u>V773 Cas</u>	56516.47530	0.00062	Prim	I	RU—this study
<u>V773 Cas</u>	56538.46837	0.00084	Sec	C	RU—this study

= S =

Table 2. Sources in the clean sample.

(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)	(12)
ID _{VUDS}	ID _{HST}	RA	Dec	z_{spec}	Flag	V	$f_{\lambda}(895)$	$\text{err} f_{\lambda}(895)$	$f_{\lambda}(1470)$	$\text{err} f_{\lambda}(1470)$	$EW_{\text{Ly}\alpha}$
		(deg)	(deg)			(mag)	(10^{-19})	(10^{-19})	(10^{-19})	(10^{-19})	(Å)
<u>COSMOS/CANDELS</u>											
510998698	6772	150.082313	2.261036	4.0651	9	25.83 ± 0.12	1.98	3.23	9.04	1.49	-24.99
511002138	4913	150.122450	2.237101	4.3600	9	24.90 ± 0.06	8.25	3.26	17.08	5.63	-19.18
511227001	20282	150.062182	2.423024	3.6350	3	25.57 ± 0.11	1.79	1.39	9.26	0.48	-23.84
5100998496	6868	150.205055	2.262162	3.8979	4	25.90 ± 0.15	-3.93	1.68	17.24	0.60	≥ 0
5101226001	20579	150.189291	2.427818	3.7327	3	25.71 ± 0.11	1.39	2.10	11.70	0.46	≥ 0
5101226251	20598	150.157203	2.42786	3.9888	3	25.42 ± 0.10	-4.95	1.40	17.74	0.52	≥ 0
5101233433*	16703	150.186851	2.379107	3.7403	4	25.19 ± 0.08	5.35	1.64	18.72	0.38	≥ 0
5101233724	16397	150.177453	2.375223	4.3862	4	26.51 ± 0.20	-1.73	1.55	9.55	0.56	-11.11
5101242274	11634	150.191107	2.317958	4.3771	4	25.95 ± 0.16	-0.34	0.80	19.61	0.67	≥ 0
<u>ECDIFS/CANDELS</u>											
530029038	3753	53.0792917	-27.8772595	4.4179	3	26.70 ± 0.19	1.27	2.41	14.79	0.56	≥ 0
530030313	4503	53.1132794	-27.8698754	3.5789	3	26.04 ± 0.09	3.83	3.87	10.62	0.61	≥ 0
530030325	4542	53.1090745	-27.8697555	3.7519	4	25.42 ± 0.06	-0.34	1.41	21.25	0.43	≥ 0
530032655	5955	53.0940950	-27.854974	3.7222	2	25.71 ± 0.07	-0.80	1.24	13.67	0.41	≥ 0
530036055	8312	53.2208728	-27.8334905	4.1608	3	25.36 ± 0.05	-1.29	1.21	27.90	0.85	≥ 0
530037593	9317	53.156571	-27.824343	3.5336	2	25.79 ± 0.08	-2.44	2.46	11.68	0.36	≥ 0
530047200	17081	53.0640040	-27.765834	3.5600	2	26.06 ± 0.09	6.36	14.9	16.27	1.49	≥ 0
530049753*	18915	53.2104941	-27.7502276	3.6055	4	25.81 ± 0.08	2.92	0.93	10.68	0.20	≥ 0
530049877	18722	53.0147863	-27.7517345	3.8245	3	25.75 ± 0.08	2.53	2.20	16.14	0.50	-8.43
530050023	19009	53.2048527	-27.7494405	3.6097	4	25.07 ± 0.04	-3.43	2.43	21.29	0.55	≥ 0
530051970	20286	53.1989481	-27.7379129	3.7983	4	24.91 ± 0.04	-1.20	1.87	26.33	0.79	-5.32
<u>ECDIFS/GEMS</u>											
530003871	958	53.196500	-28.036822	3.9022	3	26.23 ± 0.12	1.61	1.25	11.81	0.49	≥ 0
530004745	1087	53.204005	-28.03064	3.6447	3	26.24 ± 0.11	-3.48	2.8	12.27	0.66	-18.61

=a0=

THE ASTROPHYSICAL JOURNAL, 844:15 (11pp), 2017 July 20

Table 1
Physical Properties of the *Herschel* Clumps Detected in Our Selected Field
around I05480+2545 (see Figures 6(b) and (c))

ID	l (degree)	b (degree)	R_c (pc)	M_{clump} (M_{\odot})
1	183.417	-0.607	1.6	545
2	183.390	-0.502	1.0	135
3	183.351	-0.537	1.1	240
4†	183.355	-0.580	1.7	1875
5	183.327	-0.514	0.8	105
6	183.273	-0.584	1.6	595
7	183.242	-0.623	0.8	90
8	183.215	-0.588	0.9	125
9	183.168	-0.603	0.9	130
10	183.117	-0.689	1.5	335
11	183.141	-0.786	1.6	385
12	183.203	-0.661	0.6	40
13	183.230	-0.755	0.4	15
14	183.110	-0.599	0.5	40
15	183.028	-0.693	0.5	35

Note. Column 1 gives the IDs assigned to the clump. The table also lists positions, deconvolved effective radius (R_c), and clump mass (M_{clump}). The clump (ID No. 4) highlighted with a dagger contains a cluster of YSOs and the 6.7 GHz MME.

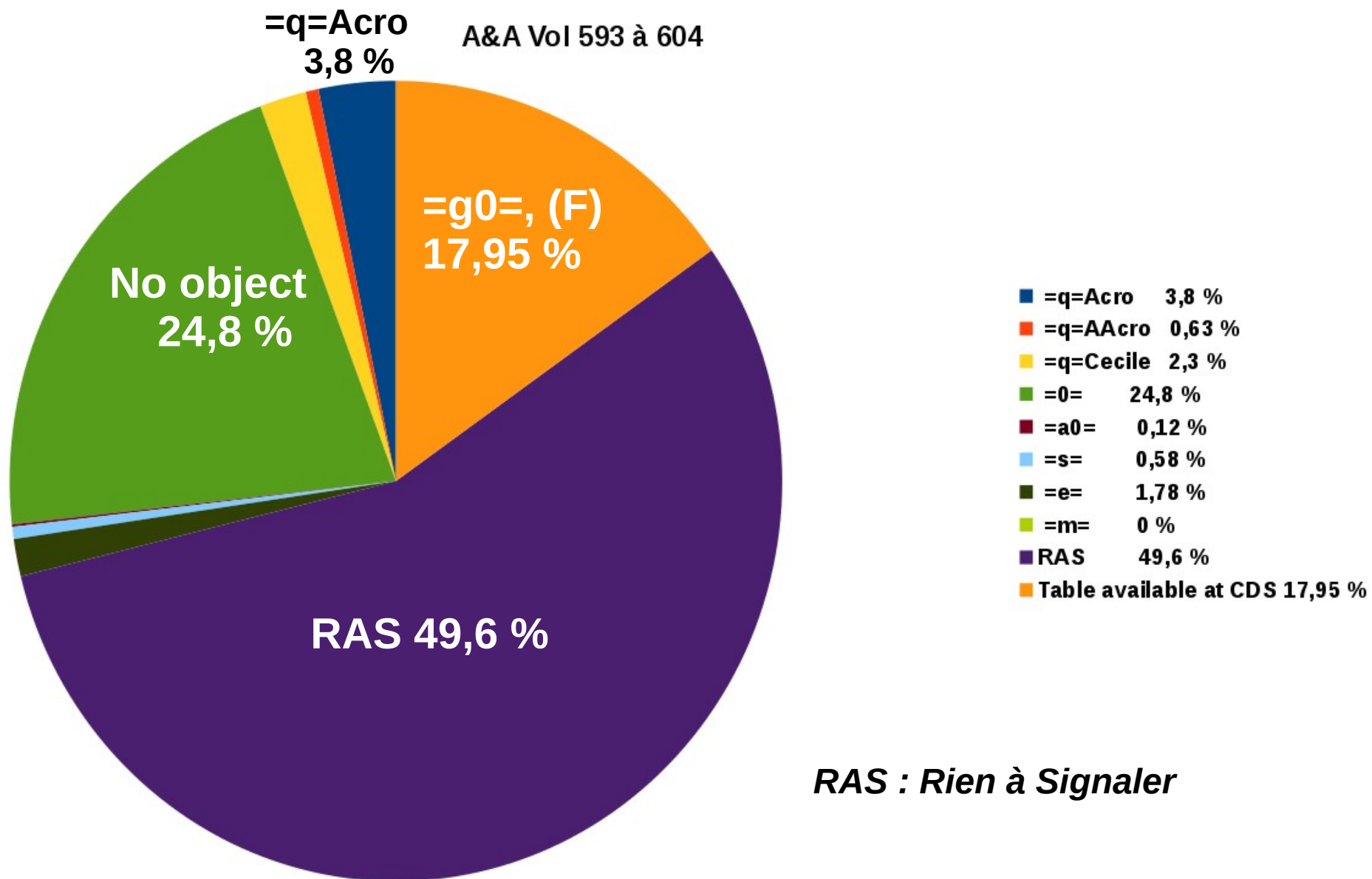
= m =

A&A 585, A71 (2016)

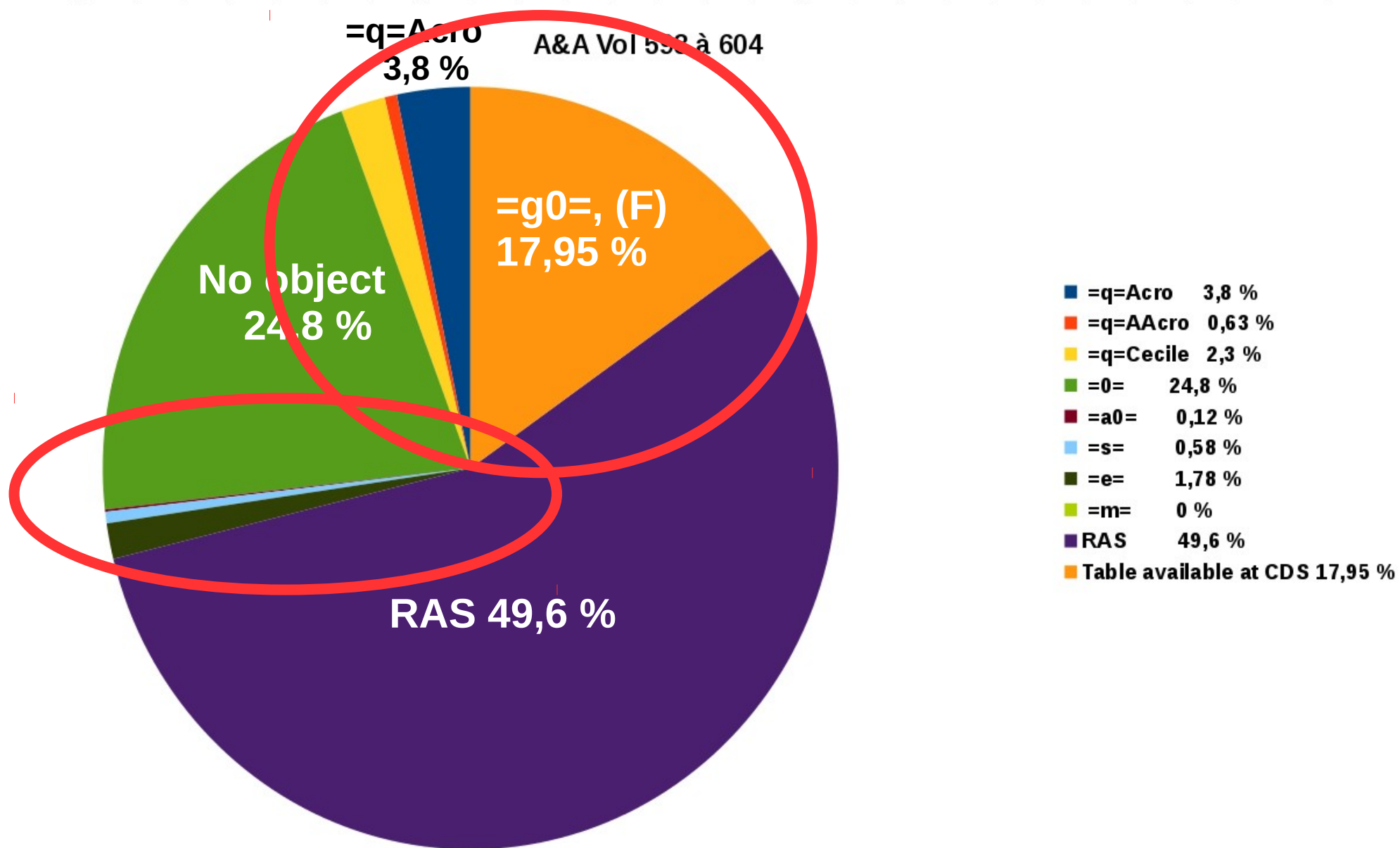
Table 1. List of BeSSeL sources observed with the JVLA.

Source	Position		V_{LSR} (km s ⁻¹)	Dist. (kpc)	L_{bol} (L_{\odot})
	RA (J2000) (h m s)	Dec (J2000) (° ' ")			
<u>G000.67-0.04</u>	17 47 20.1	-28 23 03	65.0	7.75 ± 0.72	7.3 × 10 ⁵
<u>G000.68-0.03</u>	17 47 19.9	-28 22 19	55.0	7.75 ± 0.72	3.2 × 10 ⁵
<u>G005.88-0.39</u>	18 00 30.3	-24 04 04	9	2.99 ± 0.18	5.8 × 10 ⁴
<u>G009.99-0.03</u>	18 07 50.1	-20 18 56	50	5.0 ^a	1.2 × 10 ⁴
<u>G011.92-0.61</u>	18 13 58.1	-18 54 20	30	3.37 ± 0.35	1.2 × 10 ⁴
<u>G012.43-1.12</u>	18 16 52.1	-18 41 43	-30	3.7 ^a	4.2 ^b × 10 ⁴
<u>G012.68-0.18</u>	18 13 54.7	-18 01 47	59	2.40 ± 0.18	5.7 × 10 ³
<u>G012.90-0.24</u>	18 14 34.4	-17 51 52	36	2.45 ± 0.15	8.6 × 10 ²
<u>G012.91-0.26</u>	18 14 39.4	-17 52 06	41	2.53 ± 0.20	2.7 × 10 ⁴
<u>G014.64-0.58</u>	18 19 15.5	-16 29 45	19	1.83 ± 0.07	1.1 × 10 ³
<u>G016.58-0.05</u>	18 21 09.0	-14 31 49	64	3.58 ± 0.30	1.3 × 10 ⁴
<u>G024.49-0.04</u>	18 36 05.7	-07 31 19	110	7.25 ^c ± 1.42	3.8 × 10 ⁴
<u>G026.42+1.69</u>	18 33 30.5	-05 01 02	55	3.1 ^a	9.0 ^b × 10 ³
<u>G031.58+0.08</u>	18 48 41.7	-01 09 59	96	4.90 ± 0.72	2.0 × 10 ⁴
<u>G035.02+0.35</u>	18 54 00.7	+02 01 19	52	2.33 ± 0.22	1.0 × 10 ⁴
<u>G045.07+0.13</u>	19 13 22.0	+10 50 53	59	8.00 ± 0.32	3.5 × 10 ⁵
<u>G048.61+0.02</u>	19 20 31.2	+13 55 25	25	10.75 ± 0.58	2.8 × 10 ⁵
<u>G049.19-0.34</u>	19 22 57.8	+14 16 10	67	5.29 ± 0.20	6.0 × 10 ³
<u>G074.04-1.71</u>	20 25 07.1	+34 49 58	5	1.59 ± 0.05	3.7 ^b × 10 ²
<u>G075.76+0.34</u>	20 21 41.1	+37 25 29	-10	3.51 ± 0.28	1.4 × 10 ⁴
<u>G076.38-0.62</u>	20 27 25.5	+37 22 48	-2	1.30 ± 0.09	1.4 ^b × 10 ⁴
<u>G079.88+1.18</u>	20 30 29.1	+41 15 54	-5	1.61 ± 0.07	8.6 × 10 ²
<u>G090.21+2.32</u>	21 02 22.7	+50 03 08	-3	0.67 ± 0.02	2.7 × 10 ¹
<u>G092.69+3.08</u>	21 09 21.7	+52 22 37	-10	1.63 ± 0.05	(4.7 ^b × 10 ³)

Quelques statistiques...



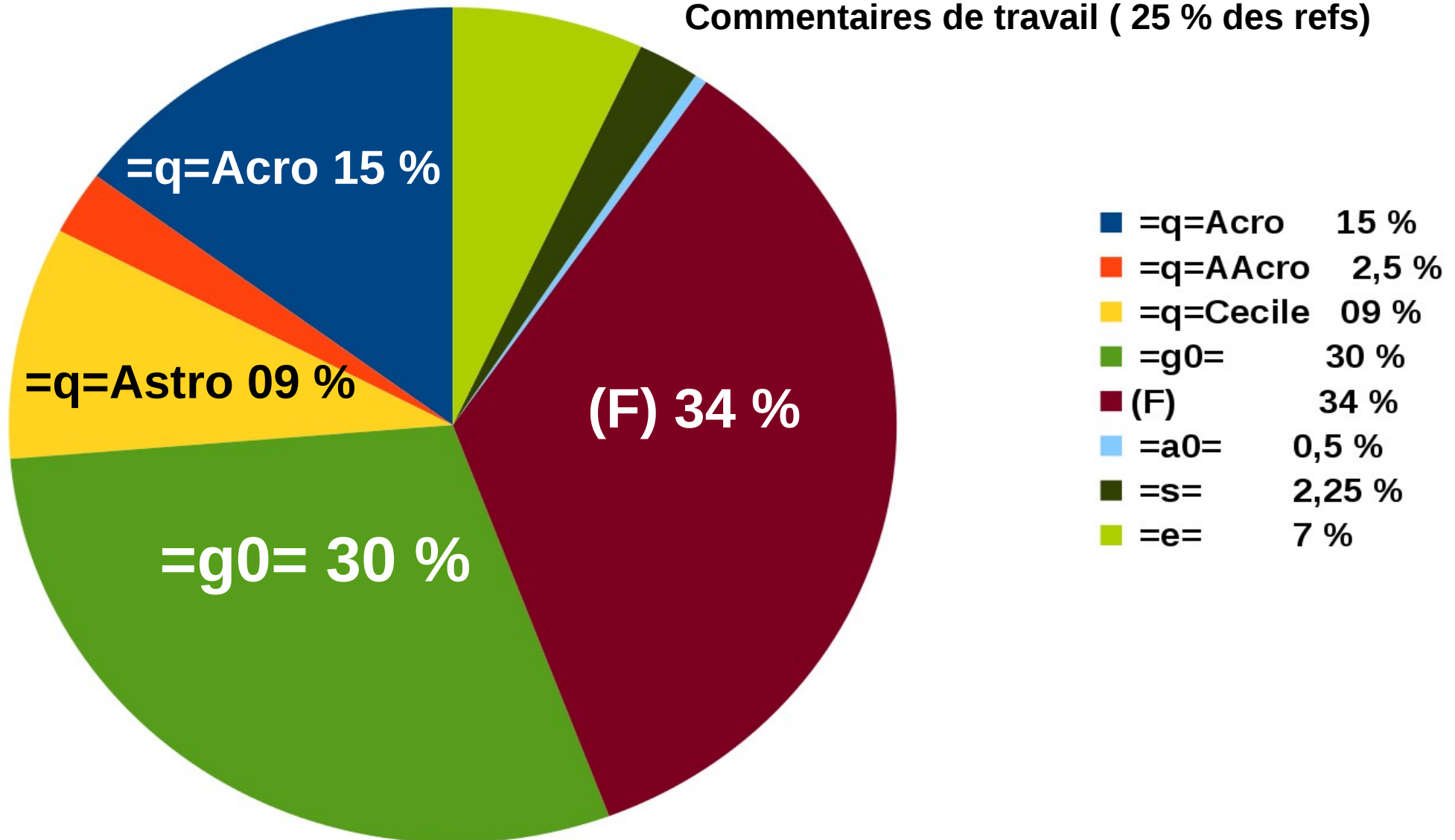
Quelques statistiques...



Quelques statistiques...

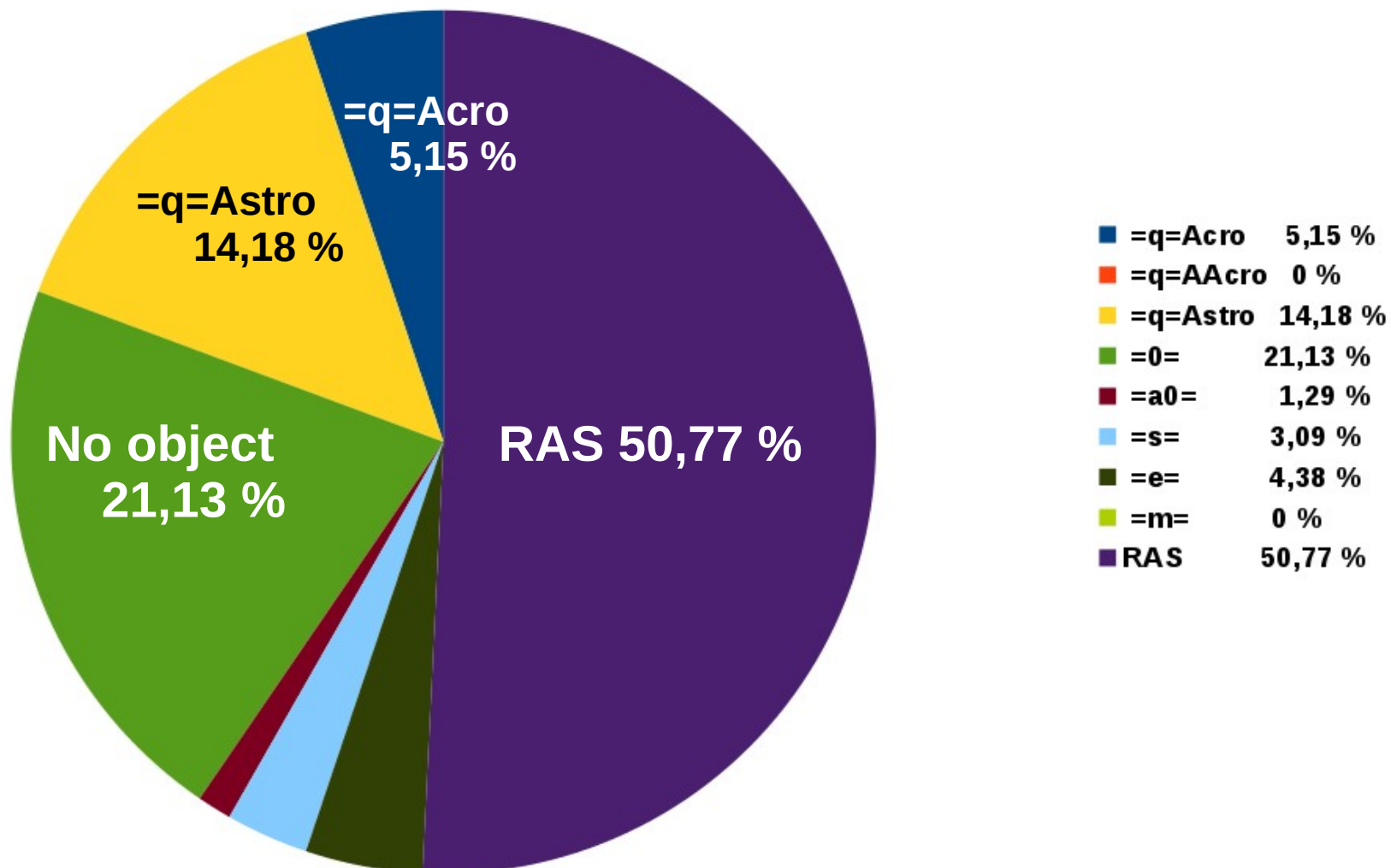
A&A

Commentaires de travail (25 % des refs)

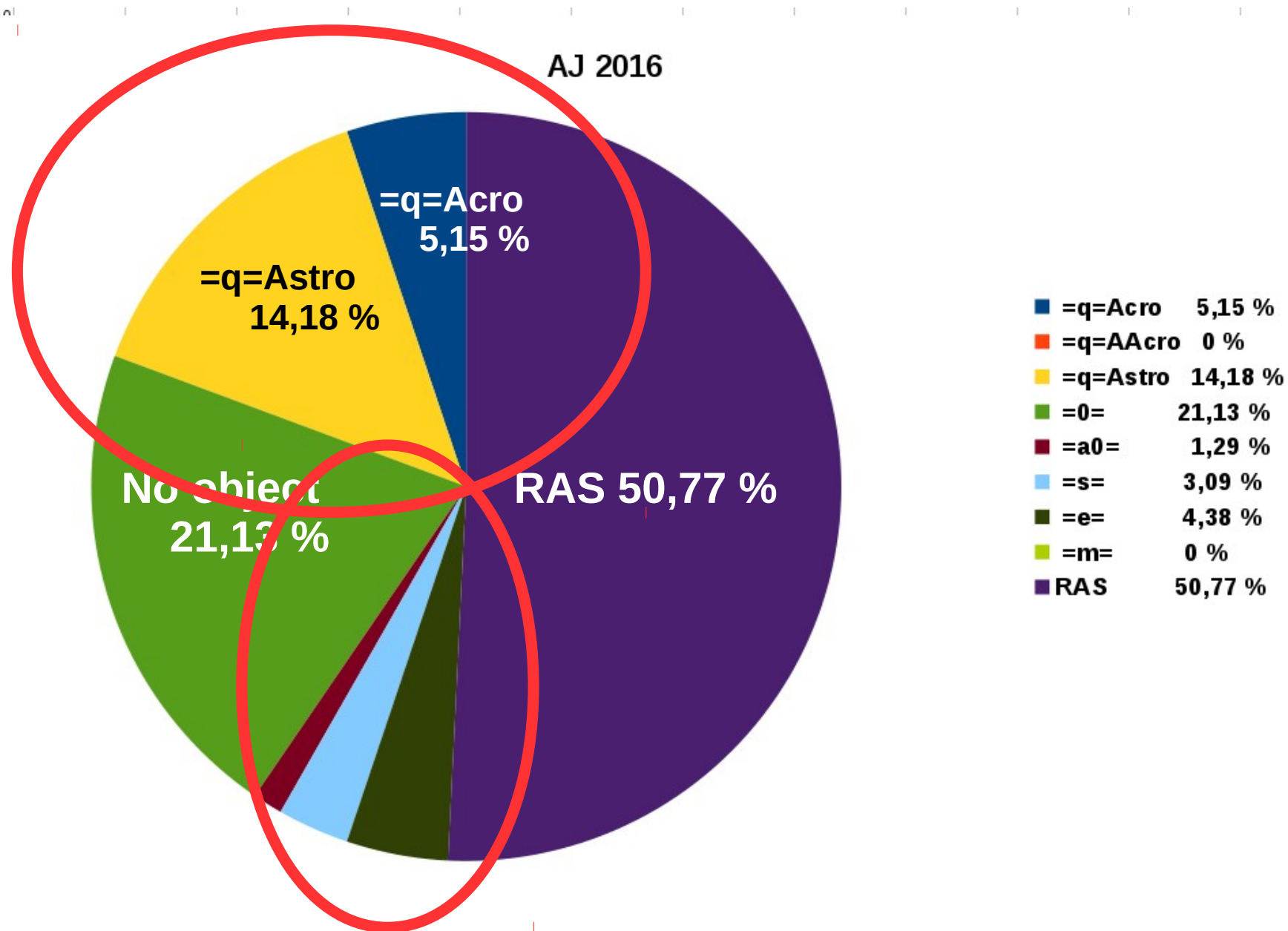


Quelques statistiques...

AJ 2016



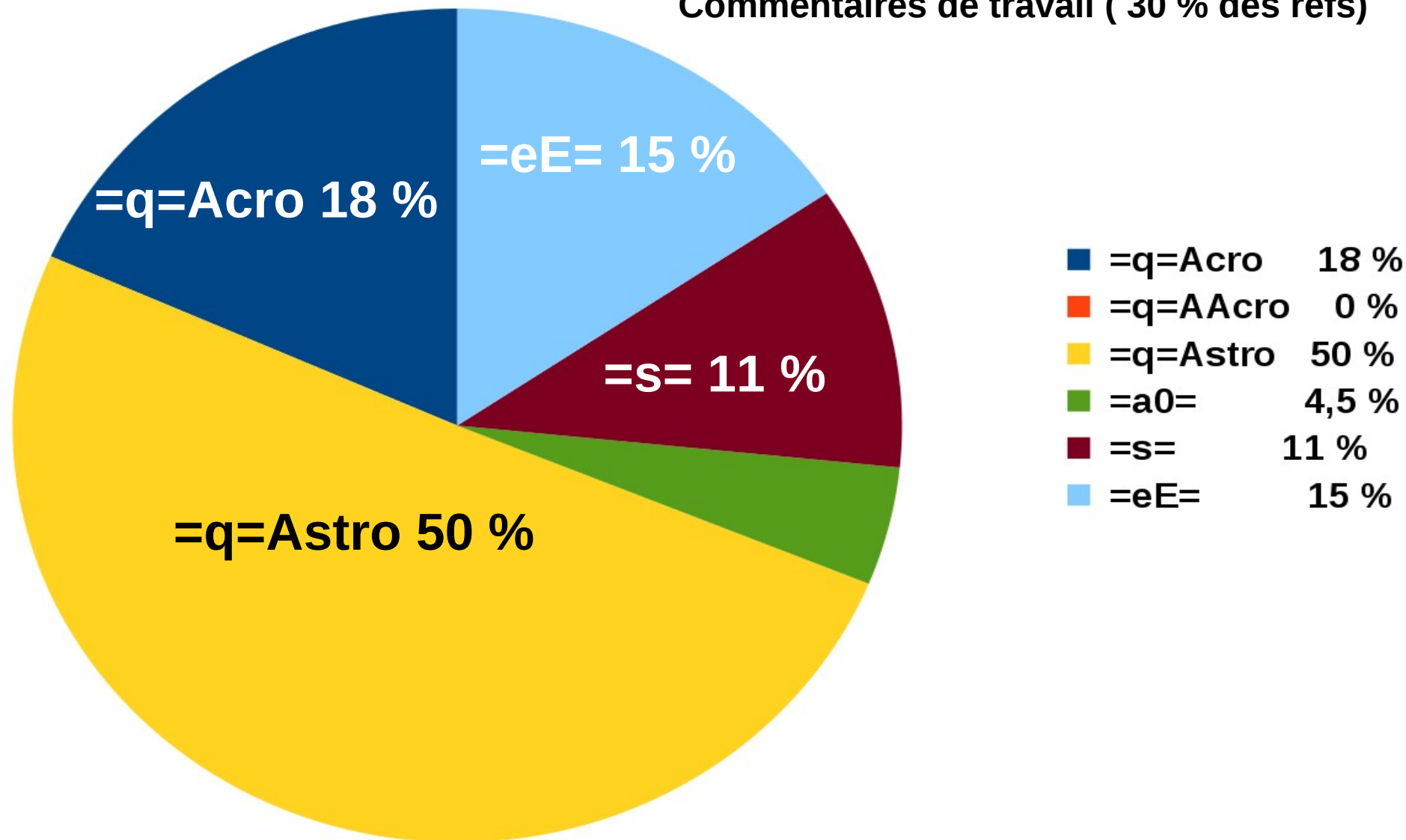
Quelques statistiques...



Quelques statistiques...

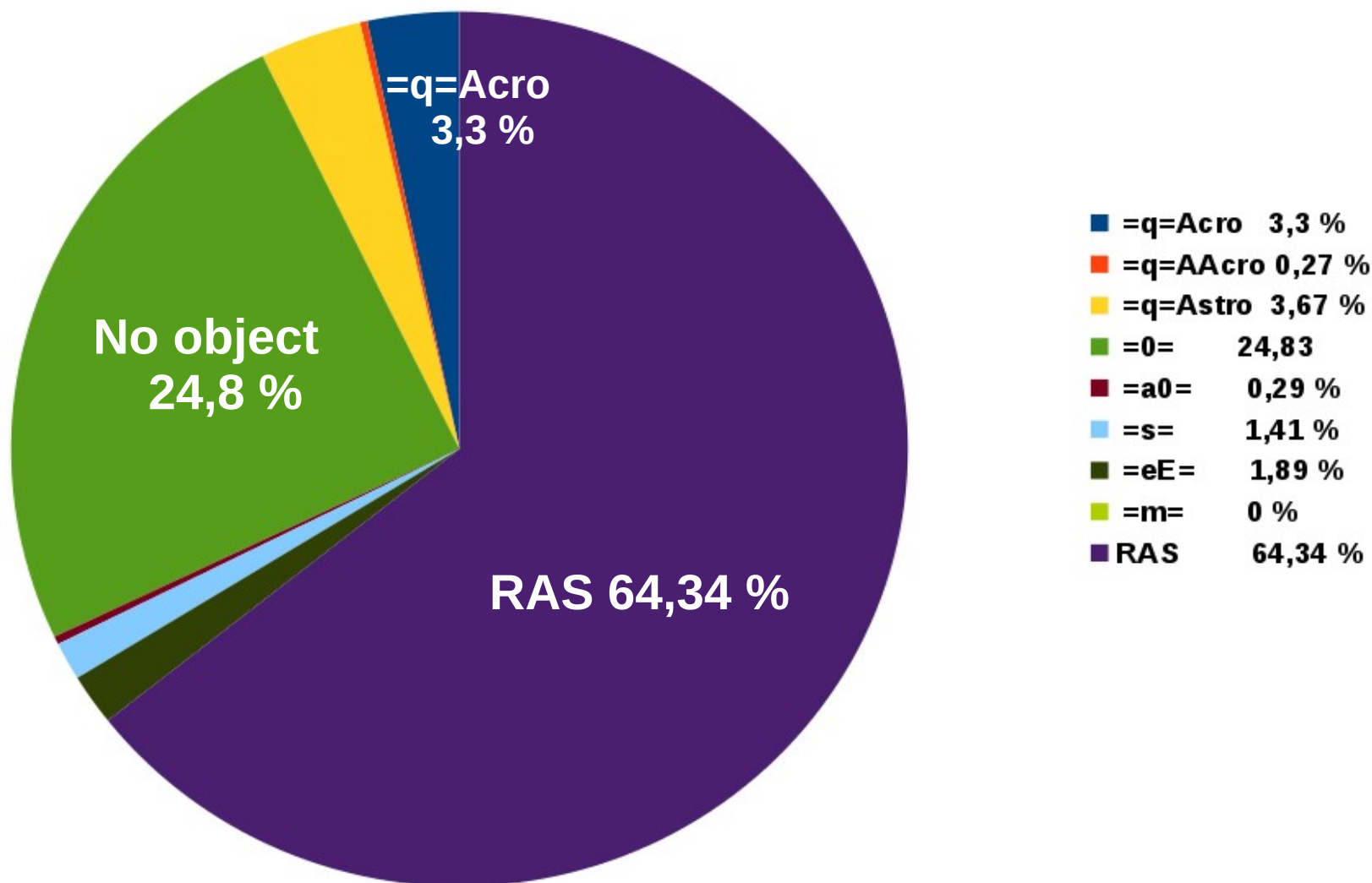
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Commentaires de travail (30 % des refs)

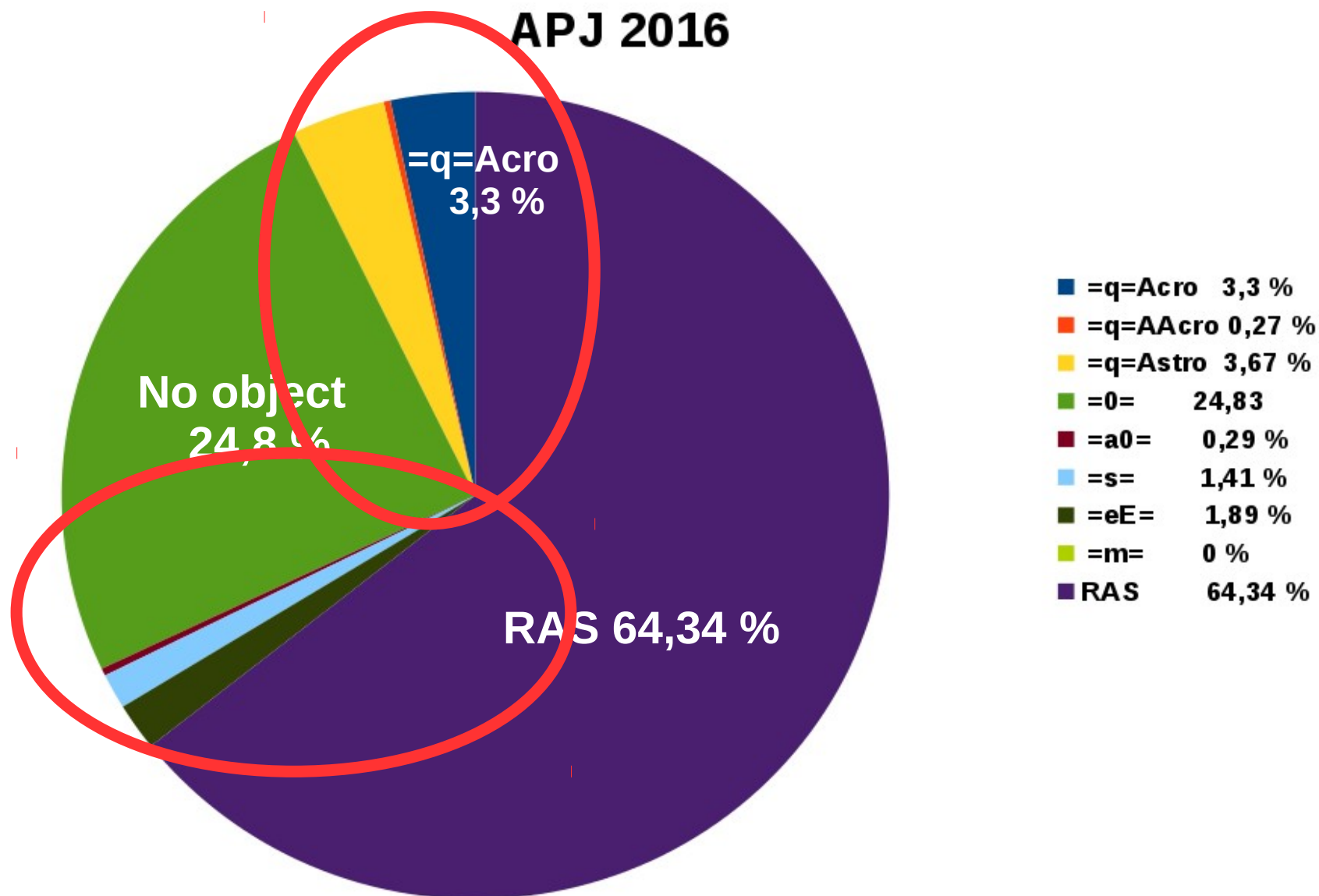


Quelques statistiques...

APJ 2016



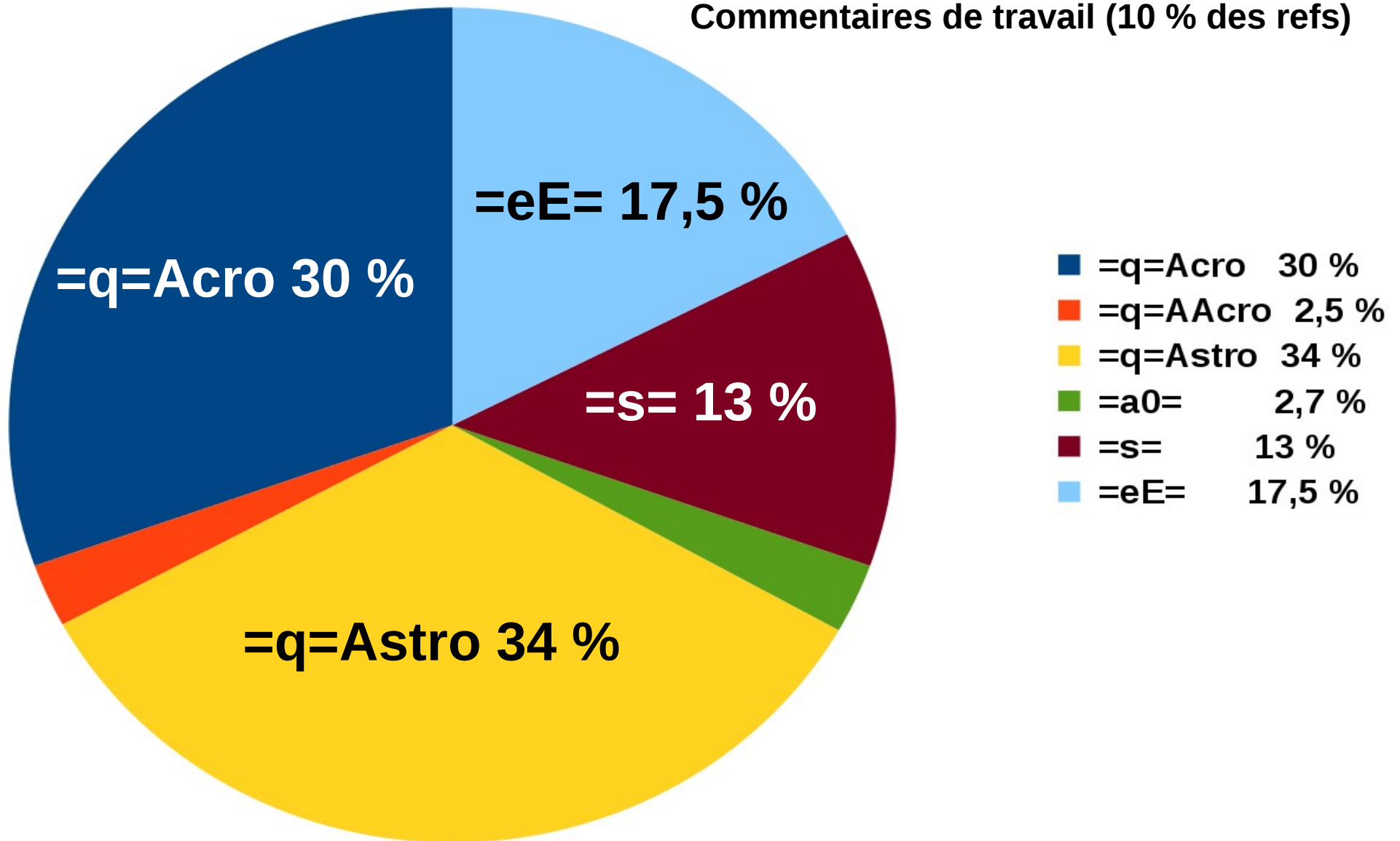
Quelques statistiques...



Quelques statistiques...

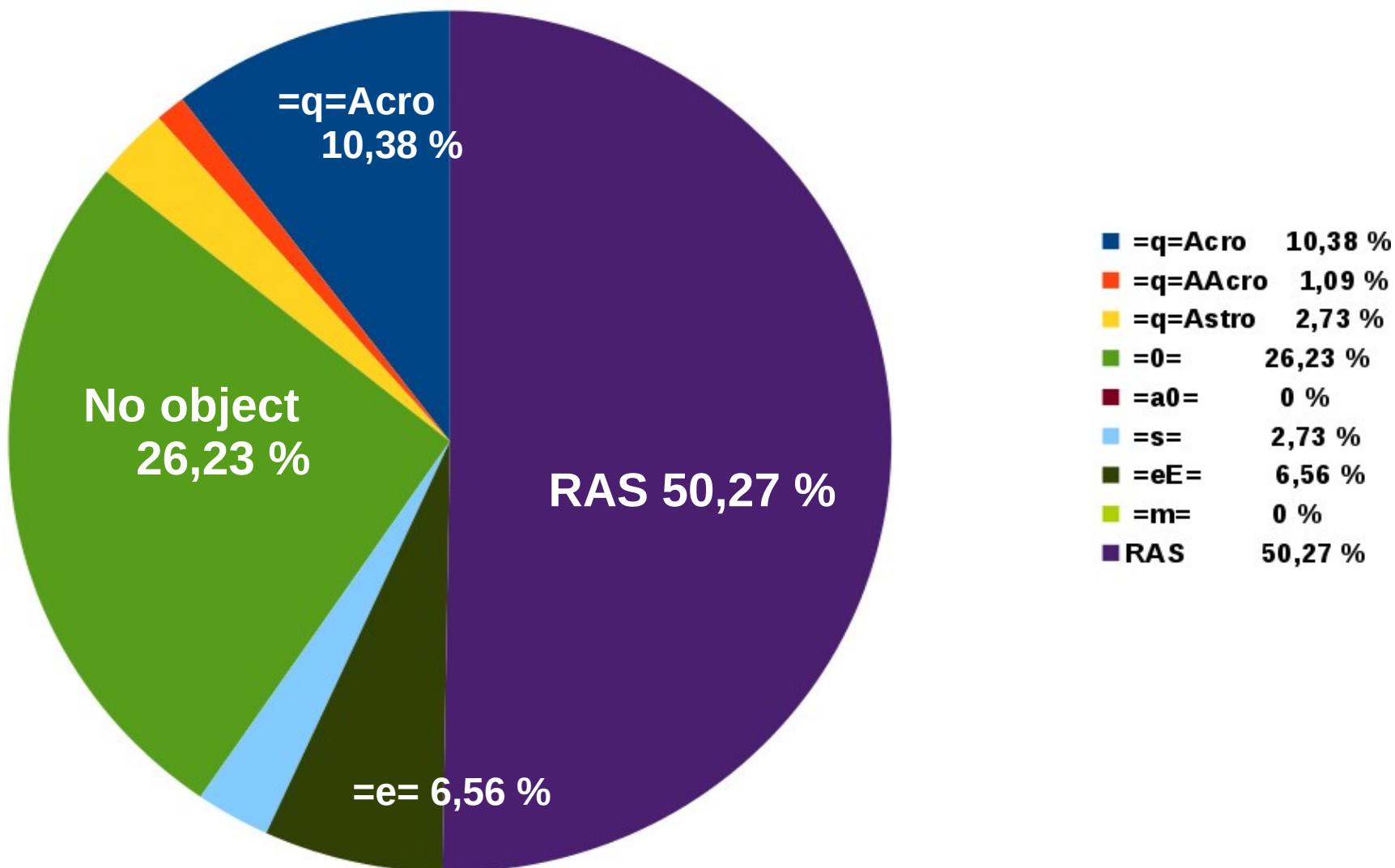
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Commentaires de travail (10 % des refs)

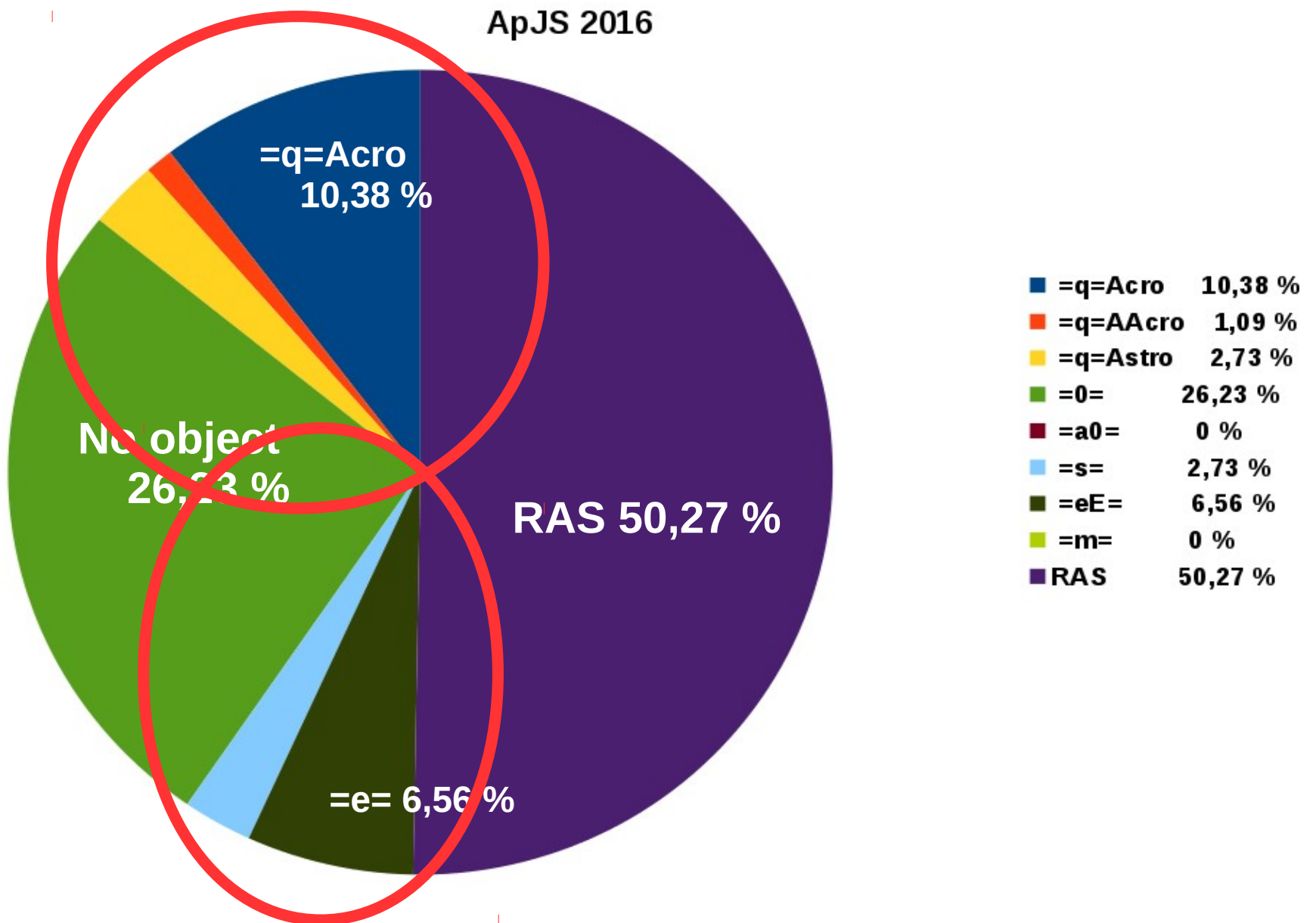


Quelques statistiques...

ApJS 2016



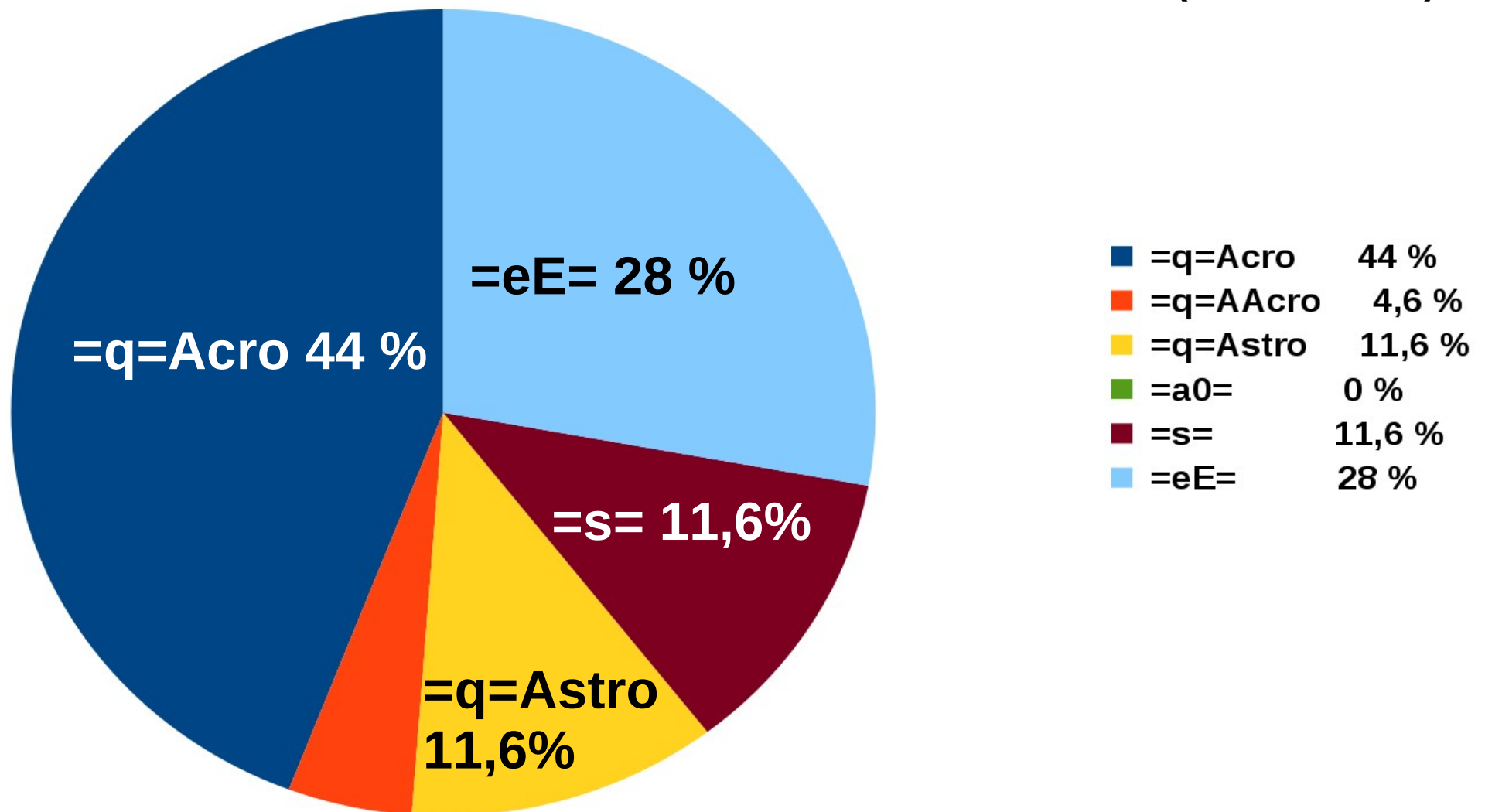
Quelques statistiques...



Quelques statistiques...

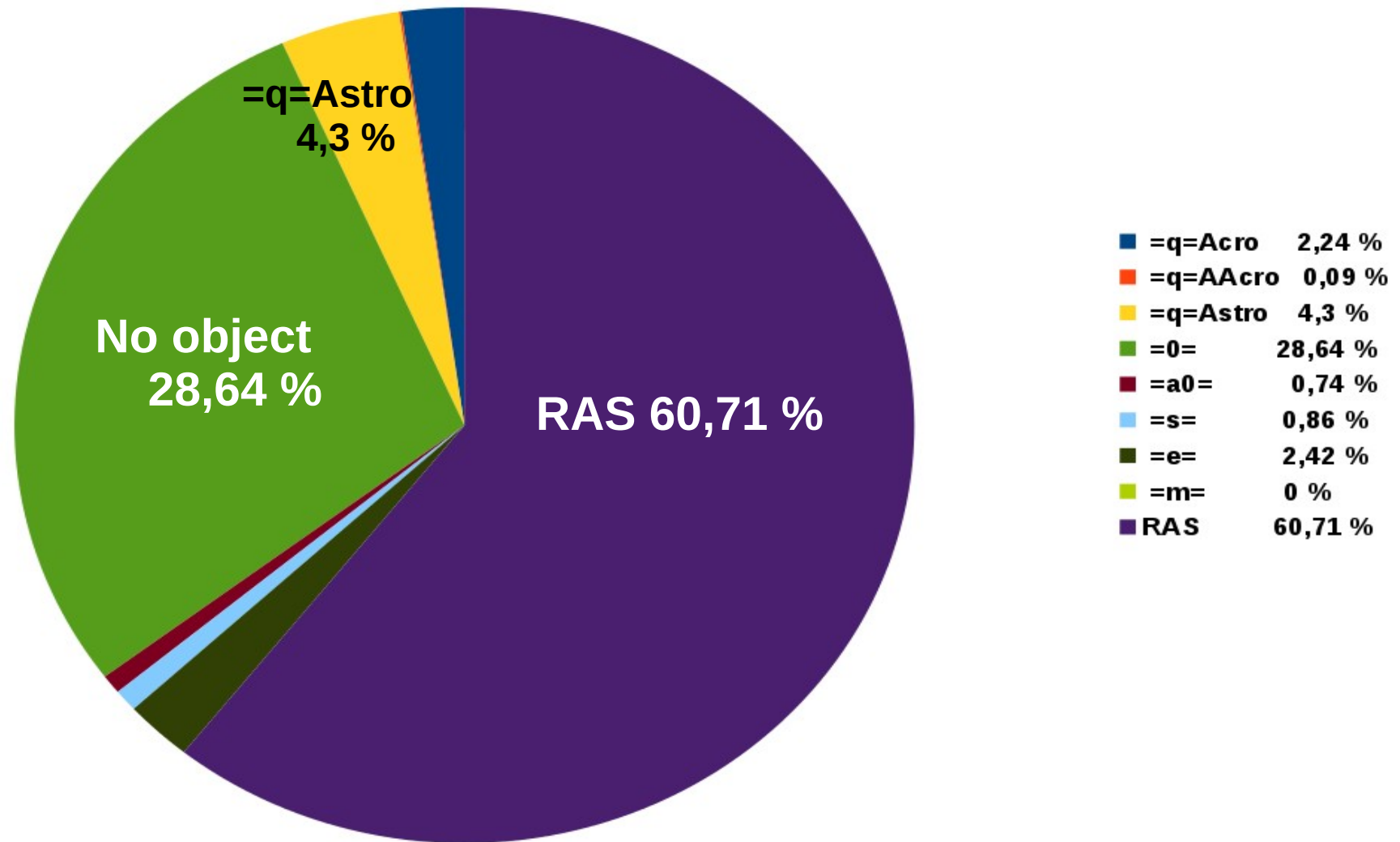
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Commentaires de travail (25 % des refs)

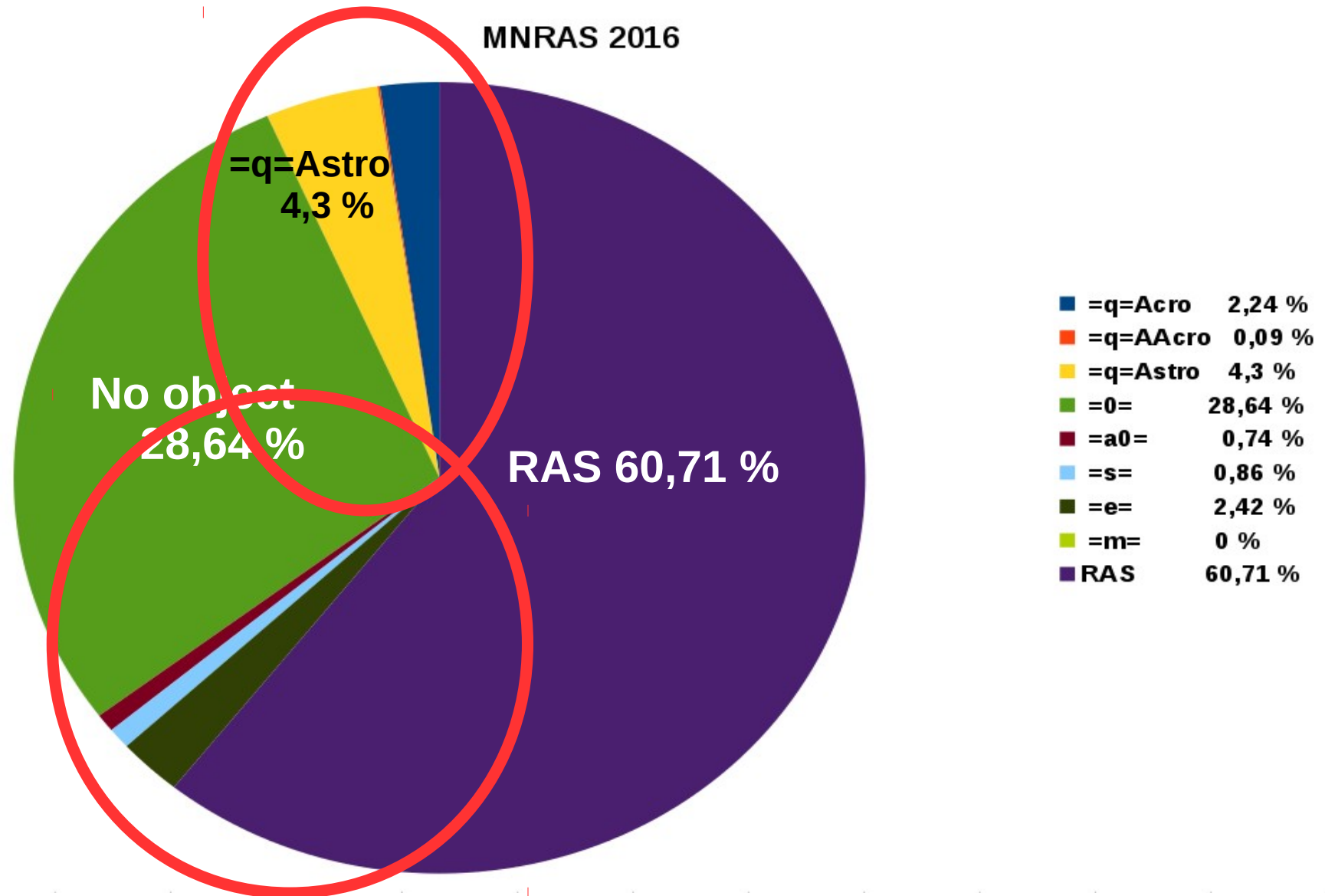


Quelques statistiques...

MNRAS 2016



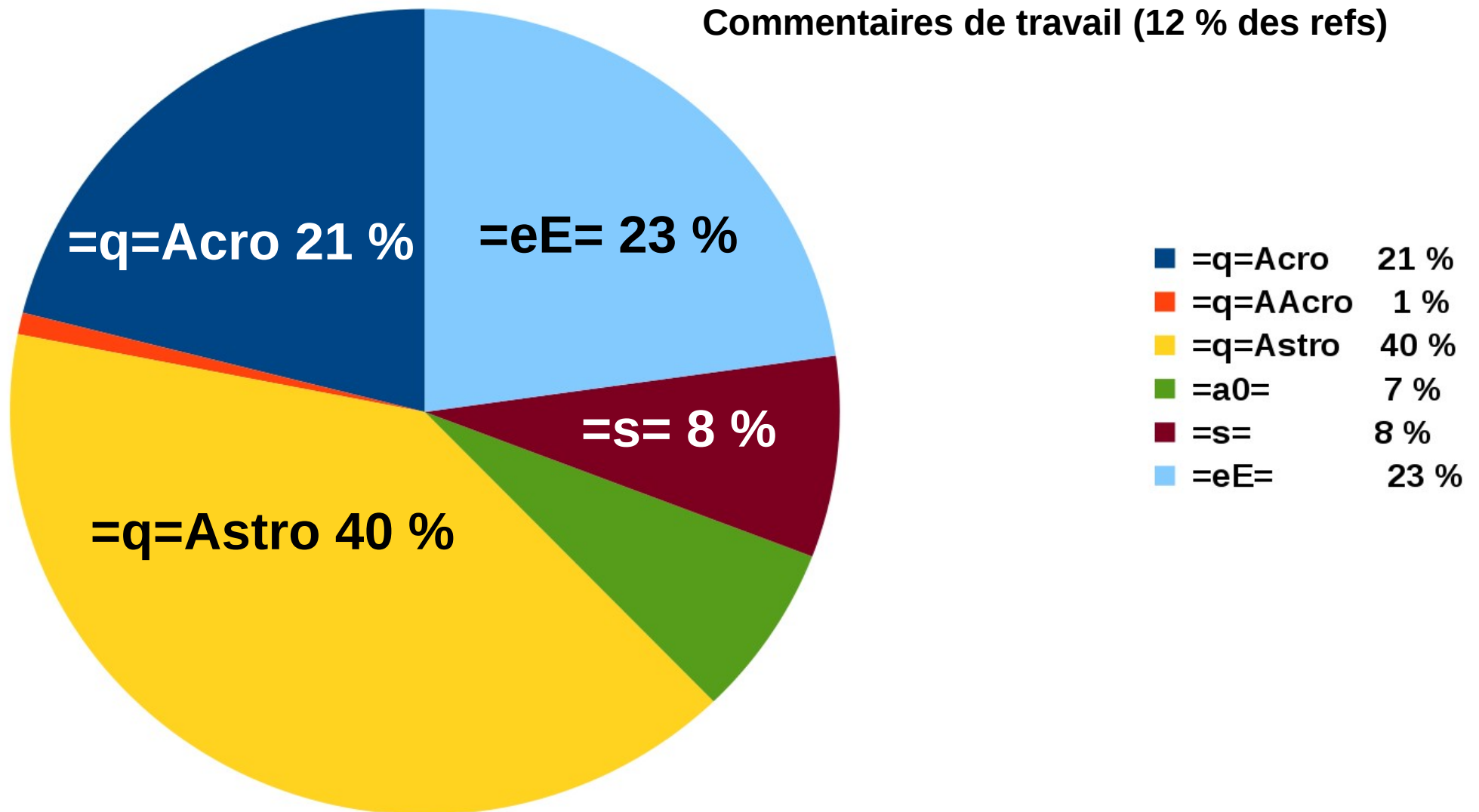
Quelques statistiques...



Quelques statistiques...

MNRAS

Commentaires de travail (12 % des refs)



Quelques statistiques...

Détails des questions à l'astronome référent

Exemple d'Evelyne (MNRAS, AJ) sur 2 ans

